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The *Herschel*^{*} PEP/HerMES Luminosity Function – I: Probing the Evolution of PACS selected Galaxies to $z \simeq 4$

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ABSTRACT

We exploit the deep and extended far-infrared data-sets (at 70, 100 and 160 μm) of the *Herschel* GTO PACS Evolutionary Probe (PEP) Survey, in combination with the HERschel Multi-tiered Extragalactic Survey (HerMES) data at 250, 350 and 500 μm , to derive the evolution of the rest-frame 35- μm , 60- μm , 90- μm , and total infrared (IR) luminosity functions (LFs) up to $z \sim 4$. We detect very strong luminosity evolution for the total IR LF ($L_{\text{IR}} \propto (1+z)^{3.55 \pm 0.10}$ up to $z \sim 2$, and $\propto (1+z)^{1.62 \pm 0.51}$ at $2 < z \lesssim 4$) combined with a density evolution ($\propto (1+z)^{-0.57 \pm 0.22}$ up to $z \sim 1$ and $\propto (1+z)^{-3.92 \pm 0.34}$ at $1 < z \lesssim 4$). In agreement with previous findings, the IR luminosity density (ρ_{IR}) increases steeply to $z \sim 1$, then flattens between $z \sim 1$ and $z \sim 3$ to decrease at $z \gtrsim 3$. Galaxies with different SEDs, masses and sSFRs evolve in very different ways and this large and deep statistical sample is the first one allowing us to separately study the different evolutionary behaviours of the individual IR populations contributing to ρ_{IR} . Galaxies occupying the well established SFR–stellar mass main sequence (MS) are found to dominate both the total IR LF and ρ_{IR} at all redshifts, with the contribution from off-MS sources (≥ 0.6 dex above MS) being nearly constant ($\sim 20\%$ of the total ρ_{IR}) and showing no significant signs of increase with increasing z over the whole $0.8 < z < 2.2$ range. Sources with mass in the range $10 \leq \log(M/M_{\odot}) \leq 11$ are found to dominate the total IR LF, with more massive galaxies prevailing at the bright end of the high- z ($\gtrsim 2$) LF. A two-fold evolutionary scheme for IR galaxies is envisaged: on the one hand, a starburst-dominated phase in which the SMBH grows and is obscured by dust, is followed by an AGN-dominated phase, then evolving toward a local elliptical. On the other hand, moderately star-forming galaxies containing a low-luminosity AGN have various properties suggesting they are good candidates for systems in a transition phase preceding the formation of steady spiral galaxies.

Key words: cosmology: observations – galaxies: active – galaxies: evolution – galaxies: luminosity function – galaxies: starburst – infrared: galaxies.

1 INTRODUCTION

Understanding the origin and growth of the galaxies we observe today is one of the main problems of current cosmology. The luminosity function (LF) provides one of the fundamental tools to probe the distribution of galaxies over cosmological time, since it allows us to assess the statistical nature of galaxy formation and evolution. When computed at different redshifts, the LF constitutes the most direct method for exploring the evolution of a galaxy population, describing the relative number of sources of different luminosities counted in representative volumes of the Universe. The LF computed for different samples of galaxies can provide a crucial comparison between the distribution of different galaxy types, i.e. galaxies at different redshifts, in different environments or selected at different wavelengths.

It has now become clear that we cannot understand galaxy evolution without accounting for the energy absorbed by dust and re-emitted at longer wavelengths (e.g. Genzel & Cesarsky 2000), in the infrared (IR) or sub-millimetre (sub-mm). Dust is responsible for obscuring the ultraviolet (UV) and optical light from galaxies: since star-formation occurs within dusty molecular clouds, far-IR and sub-mm data, where the absorbed radiation is re-emitted, are essential for providing a complete picture of the history of

star-formation through cosmic time, which is one of the fundamental instruments we have to reconstruct how galaxies have evolved since their formation epoch. For these reasons, extragalactic surveys in the rest-frame IR represent a key tool for understanding galaxy formation and evolution.

Surveys of dust emission performed with the former satellites exploring the Universe in the mid- and far-IR domain, i.e. the *Infrared Astronomical Satellite* (IRAS; Neugebauer 1984) and the *Infrared Space Observatory* (ISO; Kessler et al. 1996), allowed the first studies of the IR-galaxy LF at $z \lesssim 0.3$ (Saunders et al. 1990) and $z \lesssim 1$ (Pozzi et al. 2004), respectively. With *Spitzer* 24- μm data, it was possible to study the evolution of the mid-IR LF up to $z \sim 2$ (e.g. Le Floc’h et al. 2005, Caputi et al. 2007, Rodighiero et al. 2010a), while, even with the deepest *Spitzer Space Telescope* (Werner et al. 2004) 70- μm data, only $z \sim 1$ –1.2 could be reached in the far-IR (Magnelli et al. 2009; Patel et al. 2012) – though Magnelli et al. (2011) reached $z \sim 2$ through stacking. Since the rest-frame IR spectral energy distributions (SEDs) of star-forming galaxies and AGN peak at 60–200 μm , to measure their bolometric luminosity and evolution with z we need to observe in the far-IR/sub-mm regime. However, the detection of large numbers of high- z sources at the peak of their IR SED was not achievable before the *Herschel Space Observatory* (Pilbratt et al. 2010), due to source confusion and/or low detector sensitivity, and our knowledge of the far-IR luminosity function in the distant Universe is still affected by substantial uncertainties. Ground-based and balloon-borne observations in the mm/sub-mm range, probing the evolution of the most distant ($z \gtrsim 2$) and luminous

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dusty galaxies, have so far been limited to the identification of sources at the very bright end of the luminosity function (e.g., Chapman et al. 2005). All of these works detected strong evolution in both luminosity and/or density, indicating that IR galaxies were more luminous and/or more numerous in the past. Strong observational evidence of high rates of evolution for IR galaxies has been obtained also through the detection of a large amount of energy contained in the Cosmic Infrared Background (CIRB; Hauser & Dwek 2001), and the source counts from several deep cosmological surveys (from $15\,\mu\text{m}$ to $850\,\mu\text{m}$) largely exceeding the no-evolution expectations (e.g. Smail et al. 1997; Elbaz et al. 1999; Papovich et al. 2004; Bethermin et al. 2010; Marsden et al. 2011). Both the CIRB and the source counts require a strong increase in the IR energy density between the present time and $z\sim 1-2$. At higher redshifts the total emissivity of IR galaxies is poorly constrained, due to the scarcity of *Spitzer* galaxies at $z>2$, the large spectral extrapolations to derive the total IR luminosity from the mid-IR (see e.g. Elbaz et al. 2010, Nordon et al. 2010 and Nordon et al. 2012 for descriptions of the failure, at least at $z>1.5$, of previous total IR luminosity extrapolations from the mid-IR, although we must note that this failure mainly affects luminosity-dependent methods like, e.g., that of Chary & Elbaz 2001) and the incomplete information on the z -distribution of sub-mm sources (Chapman et al. 2005).

Herschel, with its 3.5-m mirror, is the first telescope which allows us to detect the far-IR population to high redshifts ($z\sim 4-5$) and to derive its rate of evolution through a detailed LF analysis. The new extragalactic surveys provided by *Herschel* in the far-IR/sub-mm domain, like the wide and shallow *Herschel*-ATLAS (Eales et al. 2010; Dunne et al. 2011), the complementary *Herschel* Multi-tiered Extragalactic Survey (HerMES; Oliver et al. 2012) and PACS Evolutionary Probe (PEP; Lutz et al. 2011) covering the most popular cosmological fields, and the deep, pencil beam, *Herschel*-GOODS project (Elbaz et al. 2011), will be crucial to assess galaxy and AGN evolution in the IR at $z>2$. They will give us the opportunity to study in detail the population of IR galaxies and their evolution with cosmic time since the Universe was about a billion years old. In particular, the *Photodetector Array Camera & Spectrometer* (PACS; Poglitsch et al. 2010), with its high sensitivity and resolution at $70\text{-}\mu\text{m}$, $100\text{-}\mu\text{m}$ and $160\text{-}\mu\text{m}$, is the best suited instrument to detect faint IR sources by overcoming the source confusion and blending problems that affected the previous far-IR missions.

This is the first of two papers aiming at deriving the far- and total IR LFs from the *Herschel* PACS+*Spectral and Photometric Imaging Receiver* (SPIRE; Griffin et al. 2010) data obtained within the PEP and HerMES extragalactic survey projects. In the present paper, we derive the rest-frame $35\text{-}\mu\text{m}$, $60\text{-}\mu\text{m}$, $90\text{-}\mu\text{m}$ and total IR ($8\text{--}1000\,\mu\text{m}$) LFs from a sample selected at PACS 70 , 100 and $160\,\mu\text{m}$ wavelengths in the GOODS (GOODS-S and GOODS-N), Extended Chandra Deep Field South (ECDFS) and COSMOS areas. We use the full $70\text{--}500\,\mu\text{m}$ PACS+SPIRE data to determine L_{IR} and SED properties of the PACS selected sources. In a related paper, Vaccari et al. (in prep.) derive rest-frame 100- , 160- and $250\text{-}\mu\text{m}$ and total IR LFs for a SPIRE selected sample. In addition, a third work aimed at studying the total IR LF based on the $24\text{-}\mu\text{m}$ selected sam-

ple, using all the PEP+HerMES data in the COSMOS field, is ongoing (Le Floc'h et al., in preparation).

PEP is one of the major *Herschel* Guaranteed Time extragalactic key-projects, designed specifically to determine the cosmic evolution of dusty star-formation and of the IR luminosity function. It is structured as a “wedding cake”, based on four different layers covering different areas to different depths at 100 and $160\,\mu\text{m}$ (in the GOODS-S field also at $70\,\mu\text{m}$), from the large and shallow COSMOS field to the deep, pencil beam GOODS-S field. PEP includes the most popular and widely studied extragalactic fields with extensive multi-wavelength coverage available, in particular deep optical, near-IR and *Spitzer* imaging and spectroscopic and photometric redshifts: COSMOS; Lockman Hole; Extended Groth Strip (EGS); ECDFS; GOODS-N; and GOODS-S (see Berta et al. 2010, Berta et al. 2011 and Lutz et al. 2011 for a detailed description of the fields and observations). Coordinated observations of the PEP fields at 250 , 350 and $500\,\mu\text{m}$ with SPIRE have been obtained by the HerMES Survey (Oliver et al. 2012). HerMES, analogously to PEP but extending to a much wider area, is a legacy programme designed to map a set of nested fields ($\sim 380\,\text{deg}^2$ in total) of different sizes and depths, using SPIRE (at 250 , 350 and $500\,\mu\text{m}$), and PACS (at 100 and $160\,\mu\text{m}$, shallower than PEP), with the widest component of $270\,\text{deg}^2$ with SPIRE alone. In the fields covered by PEP, the two surveys are closely coordinated to provide an optimized sampling over wavelength.

In Gruppioni et al. (2010) we started to determine the evolution with redshift of the galaxy and AGN LF in the far-IR domain by exploiting the PEP data obtained in GOODS-N by the PEP Science Demonstration Programme (SDP). Here we extend the analysis to the wider and shallower fields – COSMOS and ECDFS – and to the deepest field – GOODS-S – observed by PEP, and we also take advantage of the HerMES sub-mm data in the same fields to derive improved SED classifications and accurate total IR luminosities for our sources. This allows us to have statistically significant samples of IR galaxies at different redshifts and over a broad range of luminosities, to make a detailed study of the LF at several z intervals, all the way from $z=0$ to $z\sim 4$. The measure of the total IR luminosity obtained by integrating the SEDs, well constrained over the entire mid- and far-IR domain (and also in the sub-mm thanks to the available SPIRE data), allows us to derive the total IR LF and its evolution directly from far-IR data for unbiased samples selected at wavelengths close to the peak of dust emission. Moreover, the availability of deep multi-wavelength catalogues in the PEP fields is crucial for analysing the SEDs, obtaining k-corrections and total IR luminosities, and classifying the PEP sources into different IR populations, in order to separately study their LFs and evolutionary behaviour. This is the first study ever based on such a statistically wide and deep far-IR sample, to be able to provide LFs for different IR populations of galaxies and AGN. Here the evolution of the far- and total IR LFs (and luminosity density, hereafter ρ_{IR}) are derived up to unprecedented high redshifts (~ 4) both globally (e.g. for all the populations together) and separately, for each SED class.

Despite the abundance of information available in the literature about the stellar mass function (MF; Fontana et al. 2004; Pozzetti et al. 2010; Ilbert et al. 2010; Dominguez-

Sanchez et al. 2011), very little is known about the corresponding total IR LFs and star-formation rate (SFR) densities at different masses (an attempt based on *Spitzer* data was made by Pérez-González et al. 2005). From stellar MF studies one finds a clear increase with z of the relative fraction of massive ($\log(M/M_\odot) > 11$) star-forming objects, starting to contribute significantly to the massive-end of the MFs at $z > 1$ (Fontana et al. 2004; Ilbert et al. 2010). Their evolution and contribution to the total SFR history is however still uncertain, since only few studies have tried to reconstruct the evolution of the SF history of massive objects from optical/near-IR or mid-IR surveys (Juneau et al. 2005; Pérez-González et al. 2005; Santini et al. 2009; Fontanot et al. 2012) but none from far-IR selected surveys (providing a more direct indicator of the galaxy SF activity). In this work, we have derived the IR luminosity function and density in three different mass ranges (from $\log(M/M_\odot) = 8.5$ to $\log(M/M_\odot) = 12$), extending previous studies (limited to $z = 1.8 - 2$ for the most massive galaxies) to $z \sim 4$.

Finally, our PEP data-sets have allowed us to quantify the relative contribution of the two main modes of star formation (a relatively steady one in disk-like galaxies, defining a tight SFR-stellar mass sequence, and a starburst mode in outliers) to the total IR LF and ρ_{IR} in three redshift intervals ($0.8 < z < 1.25$, $1.25 < z < 1.8$ and $1.8 < z < 2.2$) and to test the SED-classes belonging to each mode.

The paper is structured as follows. The PEP Survey with the far-IR and multi-wavelength data, together with the SED characterisation and redshift distribution of the PEP sources, are described in Sect. 2. The LFs (rest-frame 35- μm , 60- μm , 90- μm and total IR), their evolution (derived for different SED-classes, mass and specific star-formation rate intervals) are discussed in Sect. 3. In Sect. 4 we present the number and IR luminosity densities of the different galaxy types, while in Sect. 5 we discuss our results. In Sect. 6 we present our conclusions.

Throughout this paper, we use a Chabrier initial mass function (IMF) and we assume a ΛCDM cosmology with $H_0 = 71 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_m = 0.27$, and $\Omega_\Lambda = 0.73$.

2 THE DATA

The PEP fields where we computed the LFs are: COSMOS, 2 deg² observed down to 3σ depths of $\sim 5 \text{ mJy}$ and 10.2 mJy at $100 \mu\text{m}$ and $160 \mu\text{m}$, respectively; ECDFS, $\sim 700 \text{ arcmin}^2$ down to 3σ depths of $\sim 4.5 \text{ mJy}$ and 8.5 mJy at $100 \mu\text{m}$ and $160 \mu\text{m}$, respectively; GOODS-N, $\sim 300 \text{ arcmin}^2$ to 3 and 5.7 mJy at $100 \mu\text{m}$ and $160 \mu\text{m}$, respectively; and GOODS-S, $\sim 300 \text{ arcmin}^2$ to 1.2 mJy, 1.2 mJy and 2.4 mJy at $70 \mu\text{m}$, $100 \mu\text{m}$ and $160 \mu\text{m}$, respectively. Our reference samples are the blind catalogues at 70 (in GOODS-S only), 100 and $160 \mu\text{m}$ to the 3σ level, which contain 373 (all in GOODS-S), 7176 (GOODS-S: 717, GOODS-N: 291, ECDFS: 813, COSMOS: 5355) and 7376 (GOODS-S: 867, GOODS-N: 316, ECDFS: 688, COSMOS: 5105) sources at 70, 100 and $160 \mu\text{m}$, respectively. We refer to Berta et al. (2010) and Berta et al. (2011) for a detailed description of the data catalogues and source counts.

2.1 Multi-wavelength Identification

The PEP fields benefit from an extensive multi-wavelength coverage. We have therefore associated our sources to the ancillary catalogues by means of a multi-band likelihood ratio technique (Sutherland & Saunders 1992; Ciliegi et al. 2001), starting from the longest available wavelength ($160 \mu\text{m}$, PACS) and progressively matching $100 \mu\text{m}$ (PACS), $70 \mu\text{m}$ (PACS, GOODS-S only) and $24 \mu\text{m}$ (*Spitzer*/MIPS). In the GOODS-S field, we have associated to our PEP sources the $24\text{-}\mu\text{m}$ catalogue by Magnelli et al. (2009), that we have matched with the optical+near-IR+IRAC MUSIC catalogue of Grazian et al. (2006), revised by Santini et al. (2009), which includes spectroscopic and photometric redshifts. To maximise the fraction of identifications, we limited our study to the area covered by the MUSIC catalogue ($\sim 196 \text{ arcmin}^2$), obtaining 233, 468 and 492 sources at 70, 100 and $160 \mu\text{m}$, respectively, with flux density greater than the flux limits reported above (all with either spectroscopic or photometric redshifts). In the GOODS-N field, as described in Berta et al. (2010), Berta et al. (2011) and Gruppioni et al. (2010), a PSF-matched multi-wavelength catalogue¹ was created, including photometry from the far-UV (*GALEX*) to the mid-IR (*Spitzer*). As in GOODS-S, to maximise the identifications, we limited our study in GOODS-N to the area covered by the ACS ($\sim 150 \text{ arcmin}^2$), obtaining 176 and 186 sources with flux density greater than the flux limit at 100 and $160 \mu\text{m}$, respectively (all with redshifts). We have matched our sources in the ECDFS with the multi-wavelength Survey by Yale-Chile (MUSYC) by Cardamone et al. (2010), obtaining 687 sources at $100 \mu\text{m}$ and 625 sources at $160 \mu\text{m}$ (578 and 547 with redshifts, $\sim 45\%$ spectroscopic). Finally, in COSMOS, we have matched our catalogue with the deep $24\text{-}\mu\text{m}$ sample of Le Floc'h et al. (2009) and with the IRAC-based catalogue of Ilbert et al. (2010), including optical and near-IR photometry and photometric redshifts. After the removal of PEP sources within flagged areas of the optical and/or IRAC COSMOS catalogues, we ended up with two catalogues consisting of 4110 and 4118 sources, with flux densities ≥ 5.0 and $\geq 10.2 \text{ mJy}$ at 100 and $160 \mu\text{m}$ respectively (3817 and 3849 with either spectroscopic or photometric redshifts). Throughout this paper and specifically for the SED fits described in Section 2.2, we adopt these spectroscopic or rest-frame UV to near-IR photometric redshifts for the various fields.

The HerMES extragalactic survey (Oliver et al. 2012) performed coordinated observations with SPIRE at 250, 350 and $500 \mu\text{m}$ in the same fields covered by PEP. In particular, in HerMES a prior source extraction was performed using the method presented in Roseboom et al. (2011), based on MIPS- $24 \mu\text{m}$ positions. The $24\text{-}\mu\text{m}$ sources used as priors for SPIRE source extraction are the same as those associated with our PEP sources through the likelihood ratio technique. We have therefore associated the HerMES sources with the PEP sources by means of the $24\text{-}\mu\text{m}$ sources matched to both samples. For most of our PEP sources (~ 87 per cent) we found a $> 3\sigma$ SPIRE counterpart in the HerMES catalogues.

¹ publicly available at

http://www.mpe.mpg.de/ir/Research/PEP/public_data_releases.php

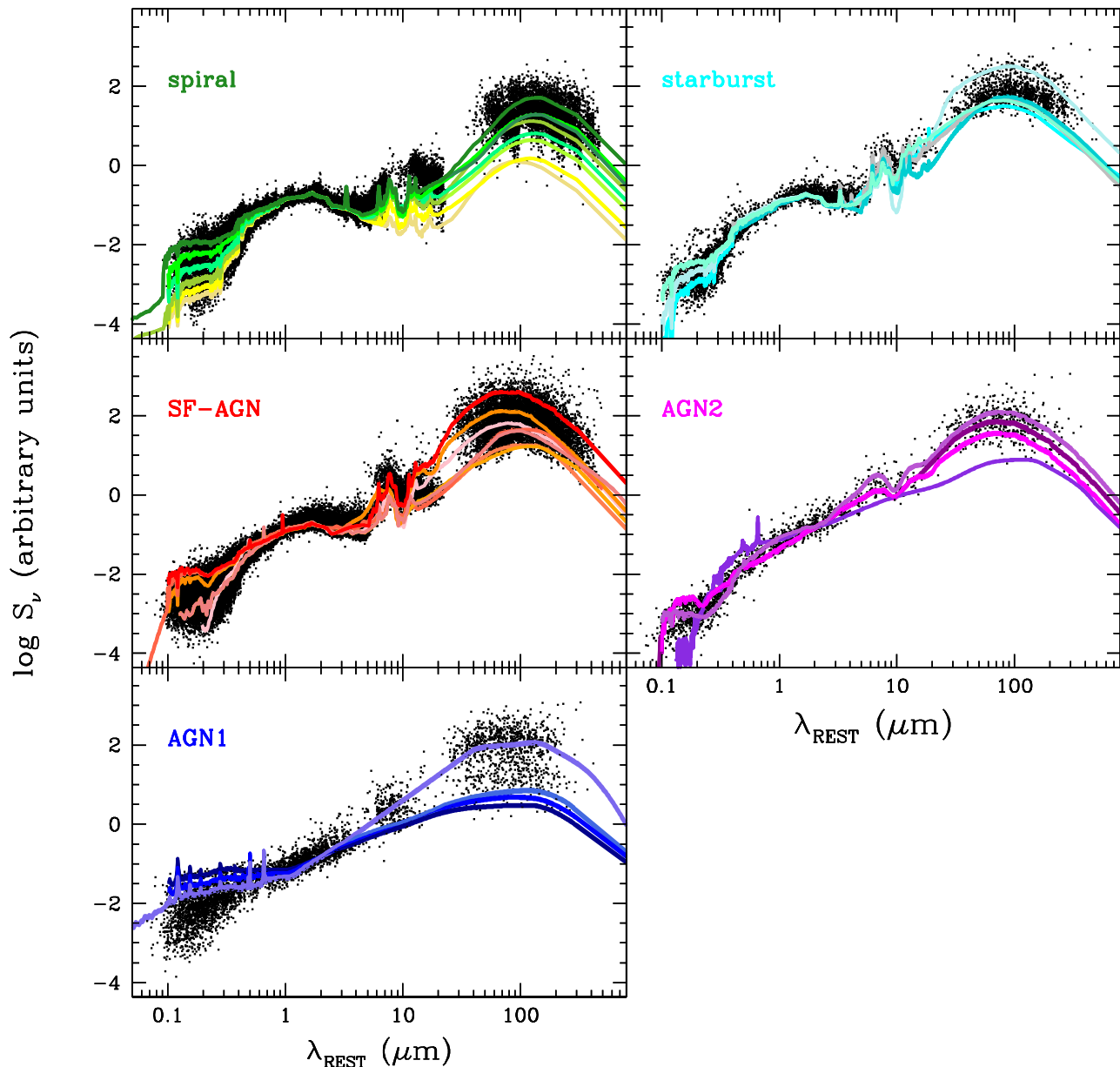


Figure 1. Observed rest-frame SEDs of the PEP sources (black dots) divided by population (as shown in the plot) and normalized to the K_s -band. The more representative templates for each SED-class have been overplotted in different colours.

2.2 Galaxy Classification

We made use of all the available multi-wavelength data to derive the SEDs of our PEP sources, which we interpreted and classified by performing a χ^2 fit (using the *Le Phare* code²; Arnouts et al. 2002 and Ilbert et al. 2006) with the semi-empirical template library of Polletta et al. (2007), representative of different classes of IR galaxies and AGN. To this library we added some templates modified in their far-IR part to better reproduce the observed *Herschel* data (see Gruppioni et al. 2010), and three starburst templates from Rieke et al. (2009). If required to improve the fit, different

extinction values ($E_{(B-V)}$ from 0.0 to 0.5) have been applied to the templates, by letting the code free to choose the most suitable extinction curve. The considered set of templates included SEDs of elliptical galaxies of different ages, lenticular, spirals (from Sa to Sdm), starburst galaxies (SB), type 1 QSOs, type 2 QSOs, Seyferts, LINERs and composite ULIRGs (containing both starburst and obscured AGN component), in the wavelength range between 0.1 and 1000 μm . The latter templates, are empirical ones created to reproduce the SEDs of the heavily obscured AGN. Two of these SEDs (the broad absorption-line QSO Markarian 231 (Berta 2005) and the Seyfert 2 galaxy IRAS 19254–7245 South (Berta et al. 2003)) are similar in shape, containing a powerful starburst component, mainly responsible for their far-IR emission, and an AGN component that contributes to

² available at <http://www.cfht.hawaii.edu/~arnouts/LEPHARE/lephare.html>

Table 1. SED Classification of the PEP Sources

field	spiral	starburst	SF-AGN	AGN2	AGN1	SF-AGN(SB)	SF-AGN(Spiral)
GOODS-S	N	N	N	N	N	N	N
70 μm	53 (23%)	22 (9%)	142 (61%)	5 (2%)	12 (5%)	26	116
100 μm	117 (25%)	60 (13%)	250 (53%)	10 (2%)	31 (7%)	54	96
160 μm	123 (25%)	55 (11%)	277 (56%)	11 (2%)	26 (6%)	73	204
GOODS-N	N	N	N	N	N	N	N
100 μm	68 (39%)	20 (11%)	78 (44%)	7 (4%)	3 (2%)	21	57
160 μm	67 (36%)	21 (11%)	85 (46%)	10 (5%)	3 (2%)	21	64
ECDFS	N	N	N	N	N	N	N
100 μm	253 (44%)	49 (9%)	245 (42%)	8 (1%)	23 (4%)	83	162
160 μm	233 (43%)	49 (9%)	231 (42%)	12 (2%)	22 (4%)	99	132
COSMOS	N	N	N	N	N	N	N
100 μm	1637 (43%)	232 (6%)	1689 (44%)	76 (2%)	183 (5%)	580	1109
160 μm	1483 (39%)	243 (6%)	1847 (48%)	103 (3%)	173 (4%)	777	1070
TOTAL	N	N	N	N	N	N	N
100 μm	2075 (41%)	361 (7%)	2262 (45%)	101 (2%)	240 (5%)	738	1424
160 μm	1906 (38%)	368 (7%)	2440 (48%)	136 (3%)	224 (4%)	970	1470

– and dominates – the mid-IR (Farrah et al. 2003), reproducing the SEDs of “obscured” AGN regardless of their optical spectra (i.e. broad or narrow lines in the optical; Gruppioni et al. 2008). Hereafter, we will refer to this class of templates and to the sources reproduced by them as to type 2 AGN (AGN2). Three other empirical templates, reproducing the observed SEDs of nearby ULIRGs containing an obscured AGN (i.e. IRAS 20551-4250; IRAS 22491-1808; NGC 6240) have been associated to the Seyfert 1.8/2, LINER ones, since they all contain an AGN, but this AGN does not dominate the observed energetic output at any wavelength (from UV to far-IR/sub-mm), showing up just in the range where the host galaxy SED has a minimum (i.e. the mid-IR). The AGN in these objects is either obscured or of low luminosity. We refer to this class as to star-forming galaxies containing an AGN (SF-AGN), since their IR luminosity is largely dominated by star-formation.

In our analysis, we make the basic assumption that the SED shapes seen at low redshifts are also able to represent the higher redshift objects. In any case, to further increase the range of SEDs in the fit, we have applied additional extinction with different extinction curves to our templates. All SED fits adopt fixed spectroscopic or photometric redshifts described in Section 2.1.

The template library used to fit our data contains a finite number of SEDs (38), representative of given classes of lo-

cal infrared objects, which do not vary with continuity from one class to another (there are large gaps in the parameter space). Therefore, the quality of the fit depends not only on the photometric errors, but also on the template SED uncertainties. For this reason, in our fitting procedure, in addition to the photometric errors on data, we need to take into account also the uncertainties due to the template SEDs discretisation and additional extinction. To do this, we have proceeded as described in detail by Gruppioni et al. (2008) and summarised as follows. First, we have run **Le Phare** on our PEP SEDs considering the nominal errors from catalogues, computing the distributions of the $(S_{\text{object}} - S_{\text{template}})_{\text{band}} / (\sigma_{\text{object}})_{\text{band}}$ values in each of the considered photometric band (where S_{object} and σ_{object} are the flux density and the relative error of the source, and S_{template} the flux density of the template in the considered band), iteratively increasing the photometric errors until we have obtained a Gaussian distribution with $\sigma \sim 1$. This corresponds to reduced χ^2 distributions peaked around 1 (as expected in the case of good fit). With the new photometric uncertainties (on average, significantly increased mainly in the optical/near-IR and SPIRE bands), we have run **Le Phare** on our sources for the second time, obtaining what we have taken as the final SED-fitting results.

The majority of our PEP sources are reproduced by templates of normal spiral galaxies (**spiral**), SB galax-

ies (**starburst**), and Seyfert2/1.8/LINERS/ULIRGs+AGN (**SF-AGN**), although different classes prevail at different redshifts and luminosities. The **spiral** SEDs show no clear signs of enhanced SF or nuclear activity (see Fig. 1), the far-IR bump being characterised by relatively cold dust ($T_{\text{dust}} \sim 20$ K). On the other hand, SB templates are characterised by warmer ($T_{\text{dust}} \sim 40\text{--}45$ K), more pronounced far-IR bumps and significant UV extinction, indicative of intense star-formation activity. Templates of star-forming galaxies containing either a low-luminosity or obscured AGN (**SF-AGN**) are characterised by a “flattening” in the $3\text{--}10\ \mu\text{m}$ spectrum (suggesting detection of an AGN in the wavelength range where the host galaxy SED has a minimum) and a far-IR bump dominated by star-formation, which is intermediate (in terms of both energy and T_{dust}) between spirals and SBs. Although they can be considered as star-forming galaxies at the wavelengths relevant to this work, we prefer to refer to them as **SF-AGN** throughout the paper, to keep in mind that they probably contain an AGN, whose presence, though not dominant in the far-IR, might be very important for analysis in other bands (e.g., in the X-rays or the mid-IR).

We note that the well-studied high- z ($\simeq 2.3$) SED of the strongly lensed sub-mm galaxy SMM J2135-0102, known as “the Cosmic Eyelash” (Ivison et al. 2010; Swinbank et al. 2010), best-fitted with our procedure by an extincted ($E(B-V) \sim 0.2$) IRAS 22491-1808 template (though rather poorly in the near-IR), does not represent the bulk of our population at high- z (> 1.5), whose SEDs are indeed well reproduced by our library of templates.

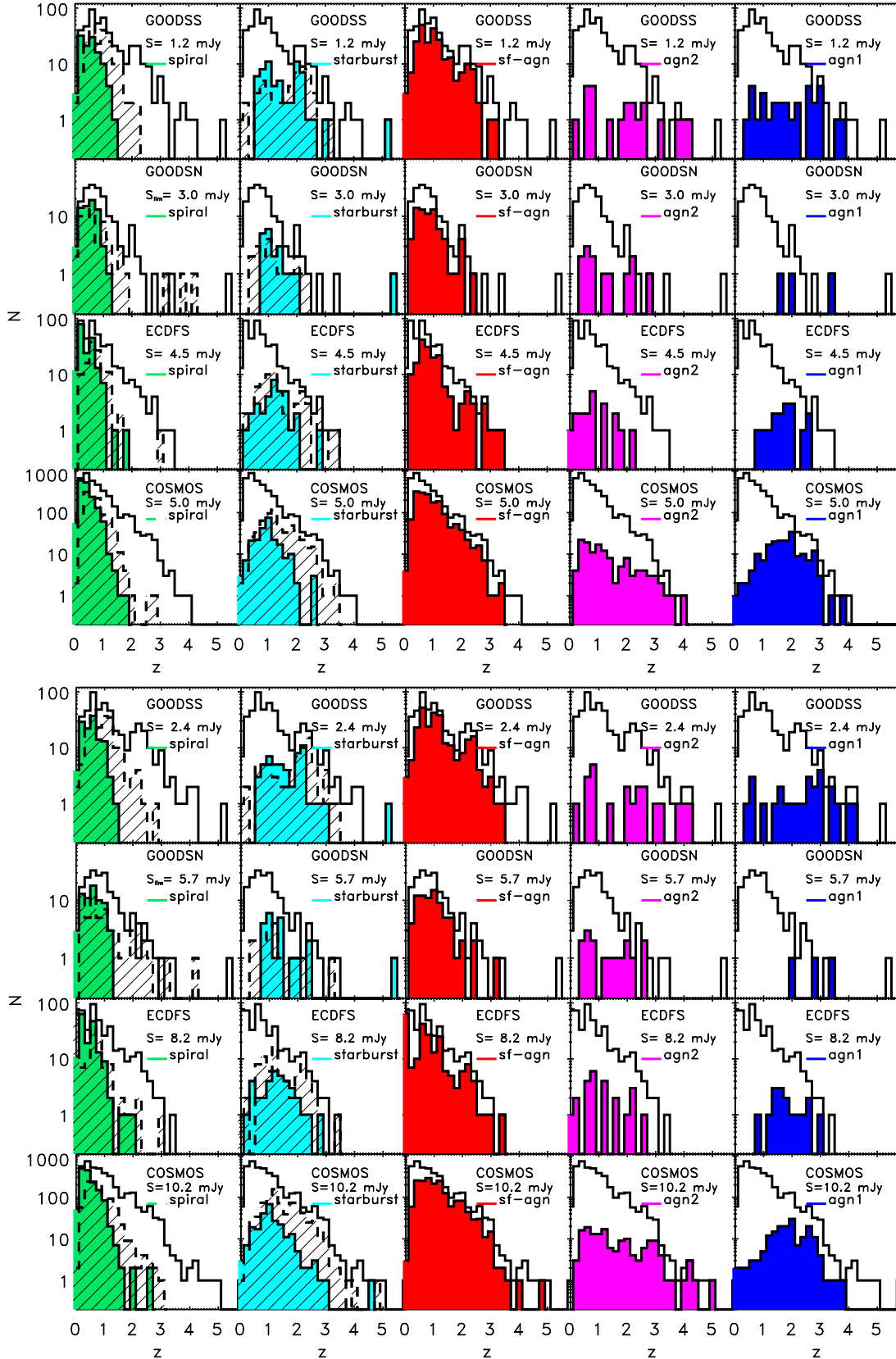
In fact, the considered template set provides very good fits to the SEDs of our PEP sources. In Fig. 1 we show the rest-frame SEDs (black dots) of the PEP sources belonging to the different “broad” SED classes (**spiral**, **starburst**, **SF-AGN**, **AGN1** and **AGN2**), compared to the template SEDs of those classes normalised to the K_s -band flux density. In Table 1 we report the fraction of sources belonging to each SED class: we find that in all the fields $\sim 41(38)$ per cent of the $100(160)\text{--}\mu\text{m}$ sources are reproduced by a **spiral** template SED, $7(7)$ per cent with a **starburst** template SED, $45(48)$ per cent with a **SF-AGN** template SED, $2(3)$ per cent with an **AGN2** SED and $5(4)$ per cent with an **AGN1** SED. We note that the fraction of **SF-AGN** derived in this work is in agreement with results from mid-IR spectroscopy (with *Spitzer-IRS*) of local star-forming galaxies from the SINGS sample by Smith et al. (2007), who found that ~ 50 per cent of local galaxies (though of lower luminosities than ours) do harbour low-luminosity AGN (of LINER or Seyfert types). Recently, Sajina et al. (2012) found an even higher fraction (~ 70 per cent) of objects hosting an AGN in the mid-IR (*Spitzer* $24\text{--}\mu\text{m}$) selected samples (~ 23 per cent AGN-dominated and ~ 47 per cent showing both AGN and starburst activity). However, since the far-IR SED of the **SF-AGN** is dominated by star-formation and at these wavelengths resembles either **starburst** or **spiral** galaxy templates, we have also divided the **SF-AGN** class into **SF-AGN(SB)** and **SF-AGN(Spiral)** sub-classes, based of their far-IR/near-IR colours (e.g. $S_{100}/S_{1.6}$) and SED resemblance (apart from the rest-frame mid-IR flattening, which is detected in all of the **SF-AGN** SEDs). Specifically, galaxies best-fitted by the Seyfert2/1.8 templates (either the original ones from Polletta et al. 2007 or those modified by Gruppioni et al. 2010) have been classified as **SF-AGN(Spiral)**, while galaxies best-fitted by the

NGC 6240, IRAS 20551-4250 or IRAS 22491-1808 templates have been classified as **SF-AGN(SB)**. The number of sources belonging to the former and the latter sub-classes are also reported in Table 1 as additional information.

2.3 Redshift Distribution

A large number of spectroscopic redshifts have been measured in the GOODS, ECDFS and COSMOS regions. In the GOODS-S and ECDFS area a collection of more than 5000 spectroscopic redshifts are available (Cristiani et al. 2000; Croom et al. 2001; Bunker et al. 2003; Dickinson et al. 2004; Stanway et al. 2004; Strolger et al. 2004; Szokoly et al. 2004; van der Wel et al. 2004; Doherty et al. 2005; Le Fèvre et al. 2005; Mignoli et al. 2005; Vanzella et al. 2008; Popesso et al. 2009; Santini et al. 2009; Balestra et al. 2010; Cooper et al. 2012). In the GOODS-N area more than 2000 spectroscopic redshifts come from various observations (Cohen et al. 2000; Wirth et al. 2004; Cowie et al. 2004; Barger et al. 2008). Finally, in COSMOS we could use a collection of ~ 3000 spectroscopic redshifts from either the public zCOSMOS bright database or the non-public zCOSMOS deep database (Lilly, S.J. et al. 2007; Lilly et al. 2009). For the PEP sources without spectroscopic redshift available, we have adopted the photometric redshifts derived from multi-wavelength (UV to near-IR) photometry by different authors in the different fields, as mentioned in Section 2.1. In the GOODS-S field the MUSIC photometric redshift catalogue (Grazian et al. 2006; Santini et al. 2009) provided photo- z s for most of our PEP sources without spectroscopic data, while in the GOODS-N field, photo- z s were obtained by Berta et al. (2010) for almost all the PEP sources within the ACS area. The Cardamone et al. (2010) and Ilbert et al. (2009) catalogues provided photometric redshifts for a large fraction of the PEP sources in the ECDFS and COSMOS areas, respectively. When considering both the spectroscopic and photometric redshifts, in our PEP fields the redshift incompleteness is very low. In particular, in the GOODS-S field we have either a spec- z or a photo- z for ~ 100 per cent of the PEP sample within the MUSIC areas (~ 80 per cent spectroscopic, though most of them lie at $z < 2.5$; see Berta et al. 2011). In the GOODS-N field we have a redshift completeness of ~ 100 per cent of sources (70 per cent spectroscopic) within the ACS area. In the ECDFS and COSMOS fields we have a redshift completeness of 88 per cent and 93 per cent respectively (45 per cent and 40 per cent spectroscopic).

The uncertainty in the photometric redshifts has been evaluated by means of a comparison with the available spec- z s by the different authors providing photo- z catalogues in the PEP fields. In particular, Berta et al. (2011) have compared the photometric and the available spectroscopic redshifts in GOODS-S, GOODS-N and COSMOS, finding a fraction of outliers, defined as objects having $\Delta z/(1+z_{\text{spec}}) > 0.2$, of ~ 2 per cent for sources with a PACS detection. Most of these outliers are sources with few photometric points available, or SEDs not well reproduced by the available templates. The median absolute deviation of the $\Delta z/(1+z_{\text{spec}})$ distribution in the three fields analysed by Berta et al. (2011) is 0.04 for the whole catalogue, and 0.038 for PACS-detected objects. In GOODS-S, Grazian et al. (2006) found an excellent agreement between photomet-



ric and spectroscopic redshifts over the fully accessible redshift range $0 < z < 6$ ($\sigma[\Delta z/(1+z)] \simeq 0.045$), with a very limited number of catastrophic errors. In COSMOS, Ilbert et al. (2010) estimated the photometric redshift uncertainties of their 3.6- μm catalogue matched with the COSMOS photo- z multi-wavelength catalogue of Ilbert et al. (2009), finding $\sigma[\Delta z/(1+z)] = 0.008$ (and < 1 per cent of catastrophic failures) at $i_{\text{AB}}^+ < 22.5$, $\sigma[\Delta z/(1+z)] = 0.011$ at $22.5 < i_{\text{AB}}^+ < 24$ and $\sigma[\Delta z/(1+z)] = 0.053$ at $24 < i_{\text{AB}}^+ < 25$. In the ECDFS, by comparing non-X-ray sources with high-quality spectroscopic redshifts, Cardamone et al. 2010 found $\sigma[\Delta z/(1+z)] = 0.008$ to $z \sim 1.2 = 0.027$ at $1.2 \leq z \leq 3.7$ and $= 0.016$ at $z > 3.7$. Note that we have checked all the $z > 2.5$ photometric redshifts through **Le Phare**, assigning the **Le Phare** derived value in case of significant disagreement with that from the catalogue (though most resulted in very good agreement). The fractions of spectroscopic redshifts in the $2.5 < z < 3.0$ interval amount to just $\sim 6\%$ and $\sim 25\%$ in COSMOS and GOODS-S respectively, dropping to $\sim 4\%$ and $\sim 6\%$ at $3.0 < z < 4.2$. We note that from our comparison between photo- and spec- z (when available) we find a general good agreement in all fields.

Photometric redshift errors may, in principle, affect the shape of the luminosity function at the bright end: by scattering objects to higher redshifts they make the steep fall-off at high luminosities appear shallower (e.g. Drory et al. 2003). To study the impact of redshift uncertainties on the inferred infrared LF, we have performed Monte Carlo simulations, as discussed in detail in Section 3.3. It is indeed very difficult to estimate the effect of catastrophic failures at $z > 2.5$, where mainly photometric redshifts are available and very little reliable spectroscopic data can be used to validate them. Moreover, for the limited high- z samples with spectroscopic information, different results are found in the different fields: i.e., in GOODS-S Berta et al. (2011) found about 25 per cent of catastrophic failures for PACS detected sources above $z \sim 2$, with the tendency to have a higher than real photometric redshift, while in COSMOS the catastrophic failures (20 per cent) found by Ilbert et al. (2009) for MIPS selected sources at $2 < z < 3$ were mostly for photo- z 's smaller than spectroscopic ones. In addition to that, sometimes high- z spectroscopic redshifts can be even more uncertain than photometric ones and great care must be taken when selecting spec- z s for comparison (i.e. we need to choose those with high quality flags). For all these reasons, we limited our analysis of photo- z uncertainties to the Monte Carlo simulations described in Section 3.3, without trying to derive uncertainties also due to catastrophic failures. In Section 3.2 we also note that the different (far-IR) photo- z approach of Lapi et al. (2011) produces consistent LF results in the common part of parameter space.

The median redshift of the 70- μm sample in GOODS-S is $z_{\text{med}}(70) = 0.67$ (the mean is $\langle z \rangle_{70} = 0.86$), while those of the 100- and 160- μm samples are different in each of the fields, given the different flux density depths reached by PEP in each area. In Table 2 we report the median and the mean redshifts found for the different fields (and for the combined sample) at the different selection wavelengths. As expected, GOODS-S reaches the highest redshifts, while the surveys in the COSMOS and ECDFS fields are shallower and sample lower redshifts. On average, the 160- μm selection favours

Table 2. Average redshifts in the PEP Survey fields

field	70 μm		100 μm		160 μm	
	z_{med}	$\langle z \rangle$	z_{med}	$\langle z \rangle$	z_{med}	$\langle z \rangle$
GOODS-S	0.67	0.86	0.85	1.07	0.98	1.16
GOODS-N			0.73	0.84	0.84	0.94
ECDFS			0.66	0.76	0.69	0.85
COSMOS			0.59	0.74	0.70	0.88
ALL FIELDS	0.67	0.86	0.64	0.78	0.73	0.91

higher redshifts than the 100- μm one (see also Berta et al. 2011).

In Fig. 2 we show the redshift distribution of the PEP sources selected at 100 μm and 160 μm in the four different fields. The black solid histogram is the total redshift distribution in the field (one for each row of the plot), while the filled histograms in different colours represent the redshift distributions of the different populations (green, **spiral**; cyan, **starburst**; red, **SF-AGN**; magenta, **AGN2**; blue, **AGN1**). The line-filled dashed histograms shown in the **spiral** and **starburst** panels represent the redshift distributions of the **SF-AGN(Spiral)** and **SF-AGN(SB)** sub-classes, respectively. In addition to the principal redshift peak, in GOODS-S a secondary peak centred at $z \sim 2$ is clearly visible at both 100 and 160 μm . A similar result has been shown and discussed also by Berta et al. (2011), while an extensive analysis of PACS GOODS-S large-scale structure at $z = 2-3$ and of a $z \sim 2.2$ filamentary overdensity have been presented by Magliocchetti et al. (2011).

3 THE LUMINOSITY FUNCTION

The sizes and depths of the PEP samples are such as to allow a direct and accurate determination of the far-IR LF in several redshift bins, from $z \simeq 0$ up to $z \sim 4$. PEP+HerMES is the unique *Herschel* survey to allow such analysis over such a wide redshift and luminosity range, sampling both the faint and bright ends of the far- and total IR LFs with sufficient statistics. Because of the redshift range covered by PEP, we would need to make significant extrapolations in wavelength when computing the rest-frame LFs at any chosen wavelengths. In order to apply the smallest extrapolations for the majority of our sources, we choose to derive the far-IR LFs at the rest-frame wavelengths corresponding to the median redshift of each sample. Given the median redshift of the 70- μm sample in GOODS-S (~ 0.67 , see Table 2), we use that sample to derive the rest-frame luminosity function at 35 μm . With the 100- and 160- μm PEP samples (whose median redshifts are ~ 0.64 and 0.73 respectively), we derived the rest-frame LFs at 60 and 90 μm . Note that, given the excellent multi-wavelength coverage available for most of our sources (thanks also to the HerMES data available in all the PEP fields and providing reliable counterparts for most of our PEP sources), their SEDs are very well determined from the UV to the sub-mm. The extrapolations are therefore well constrained by accurately defined SEDs, even at high redshifts (i.e. at $z \sim 3.5$ the rest-frame 90- μm luminosity corresponds to $\lambda_{\text{observed}} \sim 400 \mu\text{m}$, which is still in the range covered by HerMES).

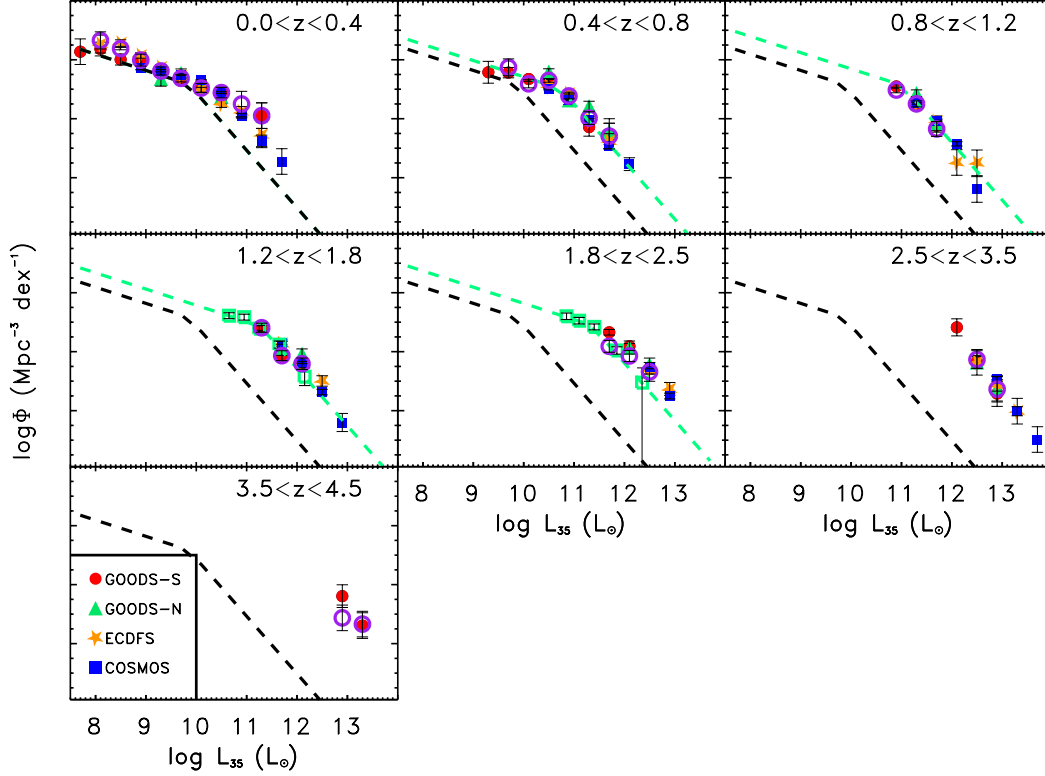


Figure 3. Rest-frame $35\,\mu\text{m}$ Luminosity Function estimated with the $1/V_{\text{max}}$ method from the PEP $70\,\mu\text{m}$ sample in the GOODS-S field (purple open circles) and independently in the four PEP fields from the $100\,\mu\text{m}$ selected samples (red filled circles, GOODS-S; green filled triangles, GOODS-N; orange filled stars, ECDFS; blue filled squares, COSMOS). The error-bars in the data points represent the Poissonian uncertainties. For comparison, we also plot the determination of Magnelli et al. (2011) at $1.2 < z < 2.5$, shown as green open squares, and the double power-law fit of Magnelli et al. (2009) and Magnelli et al. (2011), shown as a green dashed line. The black dashed line is the $z=0$ determination of Magnelli et al. (2009).

3.1 Method

The LFs are derived using the $1/V_{\text{max}}$ method (Schmidt 1968). This method is non-parametric and does not require any assumptions on the LF shape, but derives the LF directly from the data. We have first derived the LFs in each field separately, in order to check for consistency and to test the role of cosmic variance. Successively, we have made use of the whole data-sets to derive the monochromatic and total IR LFs, by means of the Avni & Bahcall (1980) method for coherent analysis of independent data-sets. We have divided the samples into different redshift bins, over the range $0 \lesssim z \lesssim 4$, selected to be almost equally populated, at least up to $z \sim 2.5$. In each redshift bin we have computed the comoving volume available to each source belonging to that bin, defined as $V_{\text{max}} = V_{z_{\text{max}}} - V_{z_{\text{min}}}$, where z_{max} is the minimum between the maximum redshift at which a source would still be included in the sample given the limiting flux of the survey (different for each field) and the upper boundary of the considered redshift bin, while z_{min} is just the lower boundary of the considered redshift bin.

When combining the four samples, we have constructed a complete sample over the whole GOODS-S+GOODS-N+ECDFS(−GOODS-S)+COSMOS region, including all the observed objects (see details in Section 3.2). The depth of the sample is not constant throughout the region, but an object with a given flux density (included in the list of

observed objects irrespective of the identity of its parent sample) can *a priori* be found in one (or more) region if its redshift is $\leq z_{\text{max}}^{\text{fld}}(S_{\text{limit}})$ of that region (e.g. sources detected in the COSMOS area are detectable over the whole joint area, while the fainter sources detected in GOODS-S are detectable in GOODS-S only). The maximum volume of space which is available to such an object to be included in the joint sample is then defined by

$$\begin{aligned}
 V_{z_{\text{max}},i} &= \frac{\Omega_{\text{GS}}}{4\pi} V_{z_{\text{max}}}^{\text{GS}} + \frac{\Omega_{\text{GN}}}{4\pi} V_{z_{\text{max}}}^{\text{GN}} + \frac{\Omega_{\text{E}}}{4\pi} V_{z_{\text{max}}}^{\text{E}} + \\
 &\quad + \frac{\Omega_{\text{C}}}{4\pi} V_{z_{\text{max}}}^{\text{C}} \\
 &\quad (\text{if } z_{\text{max},i} \leq z_{\text{max}}^{\text{C}}) \\
 &= \frac{\Omega_{\text{GS}}}{4\pi} V_{z_{\text{max}}}^{\text{GS}} + \frac{\Omega_{\text{GN}}}{4\pi} V_{z_{\text{max}}}^{\text{GN}} + \frac{\Omega_{\text{E}}}{4\pi} V_{z_{\text{max}}}^{\text{E}} \\
 &\quad (\text{if } z_{\text{max}}^{\text{C}} < z_{\text{max},i} \leq z_{\text{max}}^{\text{E}}) \\
 &= \frac{\Omega_{\text{GS}}}{4\pi} V_{z_{\text{max}}}^{\text{GS}} + \frac{\Omega_{\text{GN}}}{4\pi} V_{z_{\text{max}}}^{\text{GN}} \\
 &\quad (\text{if } z_{\text{max}}^{\text{E}} < z_{\text{max},i} \leq z_{\text{max}}^{\text{GN}}) \\
 &= \frac{\Omega_{\text{GS}}}{4\pi} V_{z_{\text{max}}}^{\text{GS}} \\
 &\quad (\text{if } z_{\text{max}}^{\text{GN}} < z_{\text{max},i})
 \end{aligned} \tag{1}$$

where $V_{z_{\text{max}}}^{\text{fld}}$ (with fld=GS, GN, E, C corresponding to GOODS-S, GOODS-N, ECDFS and COSMOS, respectively) is the comoving volume available to each source in

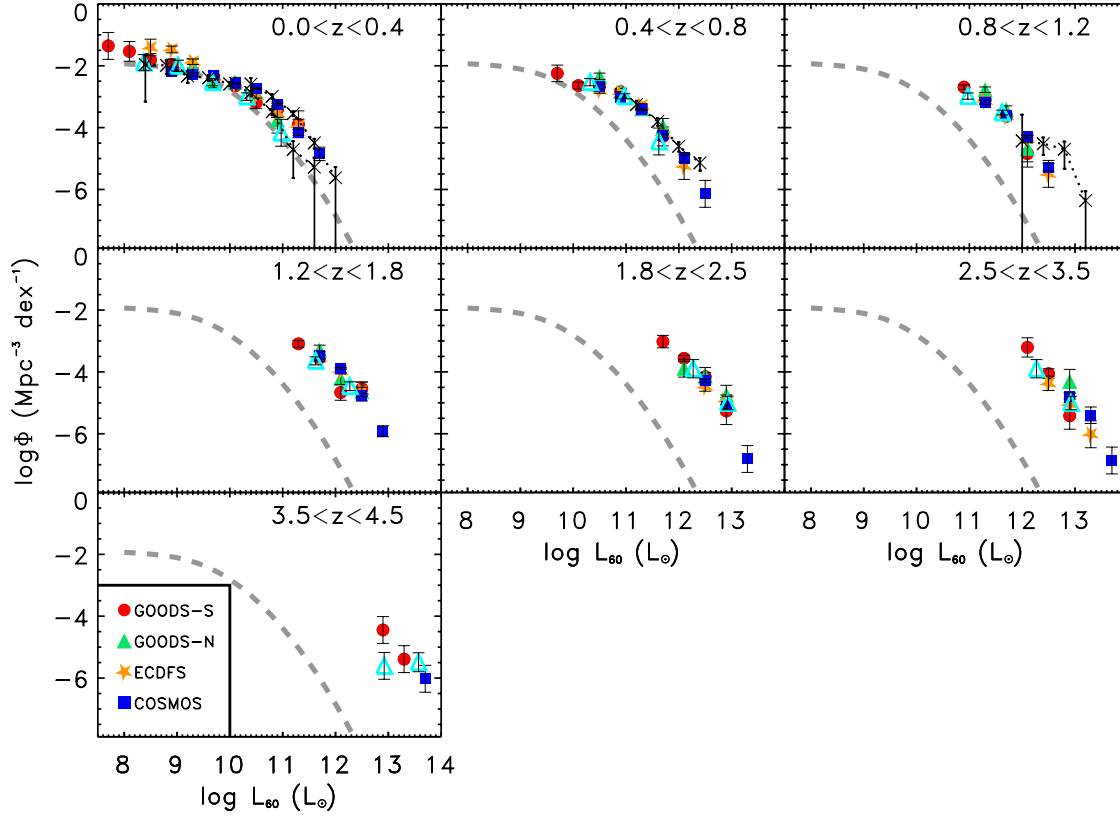


Figure 4. Rest-frame 60 μm Luminosity Function estimated independently from the 100- μm selected samples in the four PEP fields (red filled circles, GOODS-S; green filled triangles, GOODS-N; orange filled stars, ECDFS; blue filled squares, COSMOS). The diagonal crosses connected by the dotted line are the *Spitzer* 70- μm derivations of Patel et al. (2012) (in the lower z -bin we report their LFs derived both at $0.0 < z < 0.2$ and at $0.2 < z < 0.4$). The cyan open triangles show the PEP 60- μm LF from the SDP data in GOODS-N by Gruppioni et al. (2010) (the redshift bins are not exactly the same). The grey dashed line is the Saunders et al. (1990) local LF.

that field, in a given redshift bin, while Ω_{fld} is the solid angle subtended by that field sample on the sky.

For each luminosity and redshift bin, the LF is given by:

$$\Phi(L, z) = \frac{1}{\Delta L} \sum_i \frac{1}{w_i \times V_{\text{max},i}} \quad (2)$$

where $V_{\text{max},i}$ is the comoving volume over which the i -th galaxy could be observed, ΔL is the size of the luminosity bin, and w_i is the completeness correction factor of the i -th galaxy. These completeness correction factors are a combination of the completeness corrections given by Berta et al. (2010) and Berta et al. (2011), derived as described in Lutz et al. (2011), to be applied to each source as function of its flux density, together with a correction for redshift incompleteness (for the ECDFS and COSMOS only, see Section 2.3). Since, as mentioned in 2.3, the redshift incompleteness in the COSMOS and ECDFS areas is independent on PACS flux density, in these fields we have applied the corrections regardless of the source luminosity and redshift (i.e. by multiplying $\Phi(L, z)$ by 1.07 and 1.14 in COSMOS and ECDFS, respectively). However, the redshift incompleteness does not affect our conclusions, since $\gtrsim 95$ per cent of all our sources do have a redshift.

Uncertainties in the infrared LF values depend on photometric redshift uncertainties. To quantify the effects of the uncertainties in photometric redshifts on our luminosity

functions, we performed a set of Monte Carlo simulations, as described in Section 3.3.

3.2 The Rest-Frame 35-, 60- and 90- μm Luminosity Function

By following the method described above, we have derived the 35- μm , 60- μm and 90- μm rest-frame LFs from the 70- μm (in GOODS-S only), 100- μm and 160- μm samples, respectively. In order to check the consistency between the catalogues and the effects of cosmic variance, we have first derived the monochromatic LFs in each field separately. Note that, since the 70- μm data are available in the GOODS-S field only, to have a better sampling of the LF especially at the bright-end, we have also computed the rest-frame 35- μm LF from the 100- μm samples in the four fields and compared them with that obtained from the 70- μm sample (see Fig 3; Table 3). The agreement between the two derivations is very good, implying correct extrapolations in wavelength due to the good and complete SED coverage.

We have divided the samples into seven redshift bins: $0.0 < z \leq 0.4$; $0.4 < z \leq 0.8$; $0.8 < z \leq 1.2$; $1.2 < z \leq 1.8$; $1.8 < z \leq 2.5$; $2.5 < z \leq 3.5$; and $3.5 < z \leq 4.5$. The results of the computation of our 35- μm (reported in Table 3), 60- μm and 90- μm LFs are shown in Figs. 3, 4 and 5, respectively. The LFs in the four different fields appear to be consistent with each

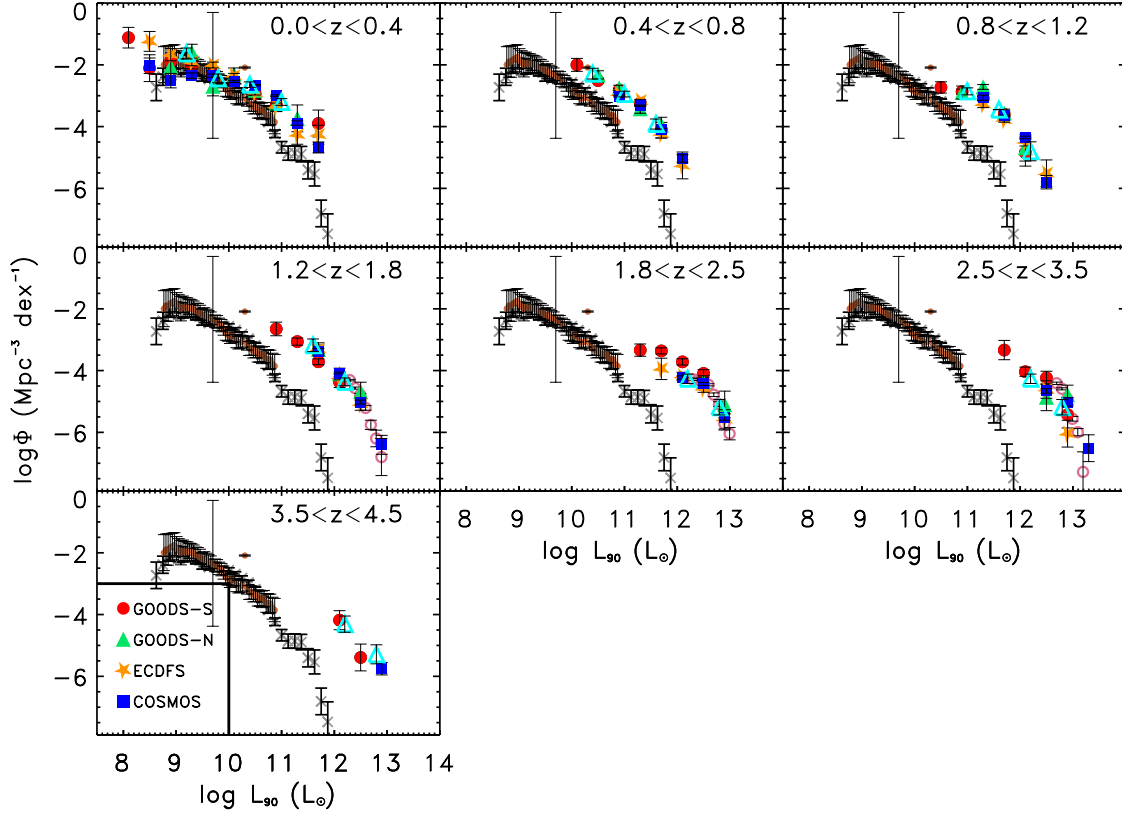


Figure 5. Rest-frame 90 μm Luminosity Function estimated independently from the 160- μm selected samples in the four PEP fields with the $1/V_{\text{max}}$ method (with the same symbols as in Fig. 4). The diagonal crosses represent the local LF of Serjeant et al. (2004), the brown dots with error bars are the local LF of Sedgwick et al. (2011) in the *AKARI* Deep Field, while the pink open circles are the 100- μm LF derivation of Lapi et al. (2011) in the H-ATLAS fields. The cyan open triangles show the PEP 90- μm LF computed by Gruppioni et al. (2010) from the SDP data in the GOODS-N (not exactly the same redshift bins).

other within the error bars ($\pm 1\sigma$) in most of the common luminosity bins. The COSMOS and GOODS-S Surveys are complementary, with the faint end of the LFs being mostly determined by data in GOODS-S, and the bright end by COSMOS data. After having checked the field-to-field consistency, we have combined the 100- and 160- μm samples in all fields, obtaining the global rest-frame 60- and 90- μm LFs (reported in Tables 4 and 5, respectively).

The data from each field in each z -bin have been plotted (and considered in the combination) only in the luminosity bins where we expect our sample to be complete, given that at fainter luminosities not all galaxy types are observable (depending on their SEDs; Ilbert et al. 2004). For comparison, we overplot the LFs at 35 μm from Magnelli et al. (2009) and Magnelli et al. (2011), the local LFs at 60 μm from Saunders et al. (1990) and those at 90 μm from Serjeant et al. (2004) and Sedgwick et al. (2011) and at 100 μm from Lapi et al. (2011), respectively. The comparison between the 35- μm LF, derived from the 70- μm PEP sample in GOODS-S, and the results of Magnelli et al. (2009) and Magnelli et al. (2011), based on a 24- μm prior extraction and stacking analysis on *Spitzer* maps, shows very good agreement, both with the data and with the double power-law fit. The $1.8 < z < 2.5$ PEP LF is consistent within $\pm 1\sigma$ with the Magnelli et al. (2011) data points, though the power-law fit at bright L_{35} ($> 10^{12} L_{\odot}$) seems to be slightly lower than our

data. At $z > 2.5$ no comparison data from *Spitzer* are available, while our LF derivation can provide hints of evolution at the bright end of the LF. In the common redshift intervals (between $z \sim 1$ and 3.5), our 90- μm LF is in very good agreement with the 100- μm Lapi et al. (2011) derivation from the H-ATLAS survey (although their redshift bins are somewhat different than ours: $1.2 < z < 1.6$; $1.6 < z < 2.0$; $2.0 < z < 2.4$; and $2.4 < z < 4.0$) and with the previous PEP-SDP derivation (Gruppioni et al. 2010). The consistency between our 90- μm LFs and the Lapi et al. (2011) ones (derived from a different sample, using a different template SED to fit the data and a rest-frame far-IR based method to derive photometric redshifts) gives us confidence that, at least up to $z \sim 3.5$, our computation is not significantly affected by incompleteness or by photo- z uncertainties.

3.3 The Total Infrared Luminosity Function

We integrate the best-fit SED of each source over $8 < \lambda_{\text{rest}} < 1000 \mu\text{m}$ to derive the total IR luminosities ($L_{\text{IR}} = L[8-1000 \mu\text{m}]$) in 11 redshift bins ($0.0-0.3$; $0.3-0.45$; $0.45-0.6$; $0.6-0.8$; $0.8-1.0$; $1.0-1.2$; $1.2-1.7$; $1.7-2.0$; $2.0-2.5$; $2.5-3.0$; and $3.0-4.2$), selected to be almost equally populated, at least up to $z \sim 2.5$. Our approach is similar to that of other studies based on mid-IR selected galaxy samples (e.g. Le Floc'h et al. 2005; Rodighiero et al. 2010a; Magnelli et

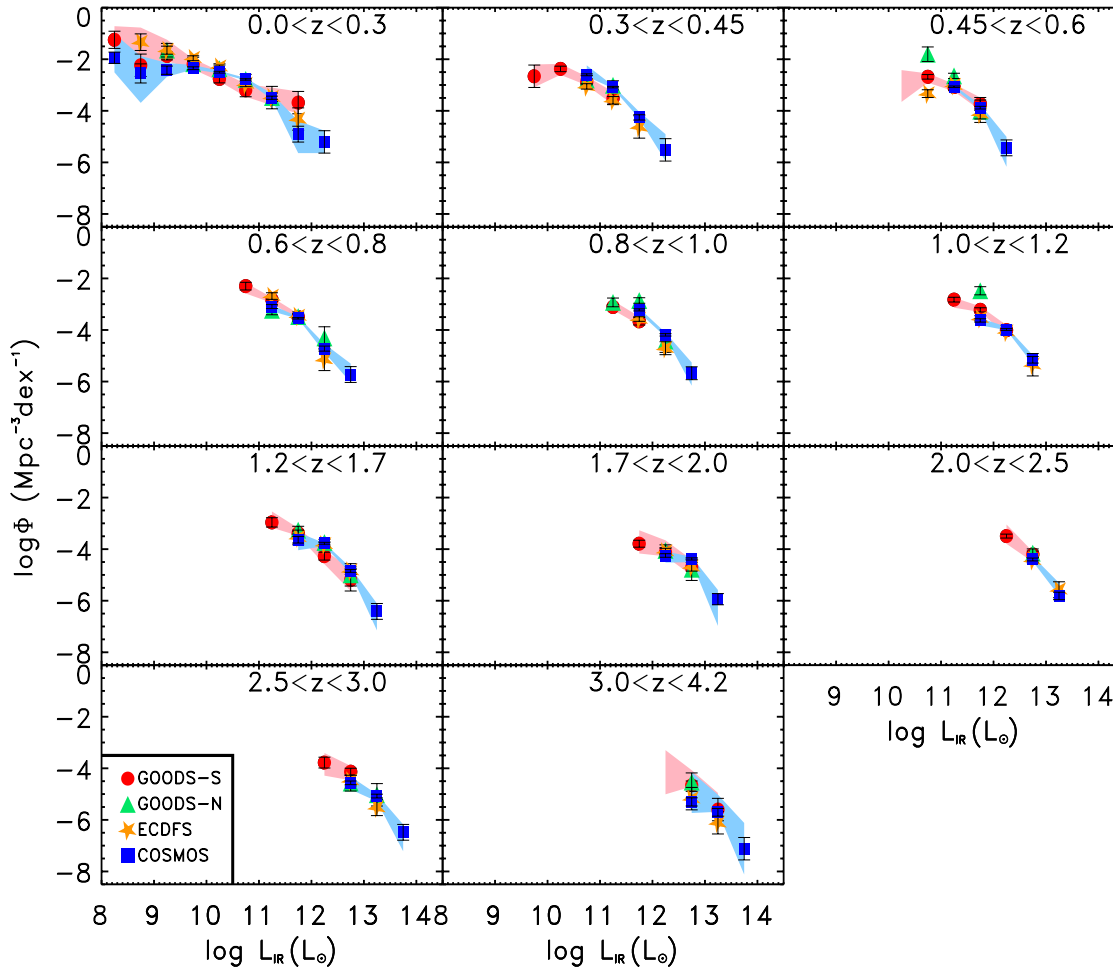


Figure 6. Total IR Luminosity Function estimated independently from the 160- μ m samples in the four PEP fields (with the same symbols as in Fig. 4). The pink and sky-blue shaded areas represent the range of values derived with 20 iterations by allowing a change in photo- z in the GOODS-S and COSMOS fields, respectively.

al. 2011), but this is the first time that the SEDs have been accurately constrained by sufficiently deep data in the far-IR/sub-mm domain and not simply extrapolated from the mid-IR to the longer wavelengths or from average flux density ratios.

As mentioned in Section 2.3, we have studied the impact of redshift uncertainties on the total IR LF by performing Monte Carlo simulations. As test cases, we used the COSMOS and GOODS-S samples (which are basically defining the bright and faint ends of the LF in an almost complementary way), and we checked the effect on the total IR luminosity function. We iterated the computation of the total IR LF by each time varying the photometric redshift of each source (assigning a randomly selected value, according to a Gaussian distribution, within the 68 per cent confidence interval). Each time, we then recomputed the monochromatic and total IR luminosities, as well as the V_{\max} value, but we did not perform the SED-fitting again, keeping the previously found best-fit template for each object (the effect on the k -correction is not relevant in the far-IR wavelength range). The results of this Monte Carlo simulation are reported in Fig. 6, where we show the total IR LFs derived in each of the PEP fields independently: the red and blue filled cir-

cles represent the estimates of the GOODS-S and COSMOS LFs, with their range of values derived with 20 iterations by allowing a change in photo- z represented by the pink and sky-blue shaded areas, respectively. The comparison shows that the effect of the uncertainty of the photometric redshifts on the error bars is slightly larger than the simple Poissonian value ($1/\sqrt{N}$), and affects mainly the lower and the higher redshift bins (especially at low and high luminosities). Using these Monte Carlo simulations, we find no evident systematic offsets caused by the photometric redshift uncertainties (see Fig. 6). This is due to the very accurate photometric redshifts available in these fields. For the total uncertainty in each luminosity bin in GOODS-S and COSMOS, we have therefore assumed the dispersion given by the Monte Carlo simulations (as shown in Fig. 7). We note that at the higher z (>2.5), where we must rely on a majority of photometric redshifts, the true uncertainties (taking into account also catastrophic failures or incompleteness effects) might be larger than derived with simulations. The unavailability of a “true” comparison sample (i.e. a large comparison sample with accurate spectroscopic redshifts and fully representative for the PACS flux selection) at high z does not allow to properly quantify this statement.

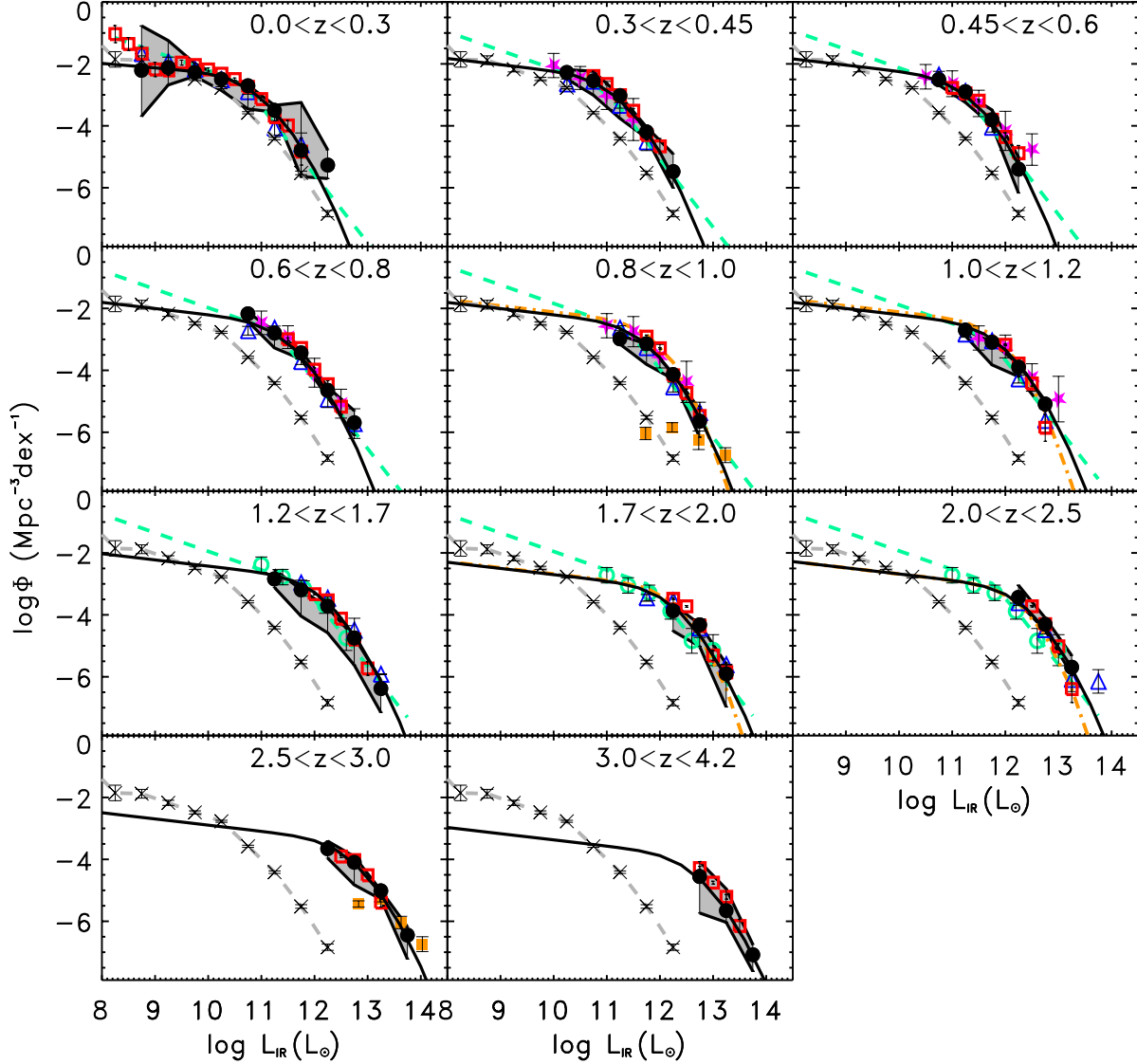


Figure 7. Total IR Luminosity Function estimated by combining the data from the four PEP fields using the Avni & Bahcall (1980) method (black filled circles). The grey filled area shows the uncertainty locus obtained by combining the Poissonian error with the photometric redshift uncertainty derived through Monte Carlo simulations. The black solid line represents our best fit to the PEP data in the different redshift bins, corresponding to the parameters reported in Table 7. Other results from the literature are plotted for comparison (diagonal crosses connected by a grey dashed line, LLF of Sanders et al. 2003; magenta filled stars, Le Floc’h et al. 2005; orange filled squares, Chapman et al. 2005; orange dot-dashed line, Caputi et al. 2007; blue open triangles, Rodighiero et al. 2010a; green dashed line, Magnelli et al. 2009, 2011; green open circles, Magnelli et al. 2011). Note that the Magnelli et al. (2011) data correspond to slightly different redshift bins: $1.3 < z < 1.8$ and $1.8 < z < 2.3$, respectively. We have therefore plotted the data points corresponding to the first redshift bin in our $1.2 < z < 1.7$ panel and the data point corresponding to the second redshift bin in both our $1.7 < z < 2.0$ and $2.0 < z < 2.5$ panels. The red open squares are the total IR LFs derived by Vaccari et al. (in preparation) from the HerMES $250\text{-}\mu\text{m}$ selected COSMOS sample.

In Fig. 7 the total IR LFs obtained by combining the $160\text{-}\mu\text{m}$ selected samples with the Avni & Bahcall (1980) technique is plotted and compared with other derivations available in the literature. The total IR LF of Sanders et al. (2003) is plotted as a local reference, in addition to the LFs of Le Floc’h et al. (2005), Rodighiero et al. (2010a), Caputi et al. (2007), Magnelli et al. (2009) and Magnelli et al. (2011) in various redshift intervals. Globally, data from surveys at different wavelengths agree relatively well over the common z -range. No data for comparison are available

at $z > 2.5$, apart from the IR LF of sub-mm galaxies from Chapman et al. (2005) and Wall et al. (2008) at $z \sim 2.5$, which represent reasonably well just the very bright end of the total IR LF. Our derivation is the first at such high redshifts, especially in the $3 < z \leq 4.2$ range. Note the good agreement between our PEP-based total IR LF and the HerMES-based one derived by Vaccari et al. (in preparation), shown by the red open squares in Fig. 7. Though consistent within the error-bars, in the highest redshift bin the HerMES LF is slightly higher than ours. Since the $250\text{-}\mu\text{m}$ HerMES selec-

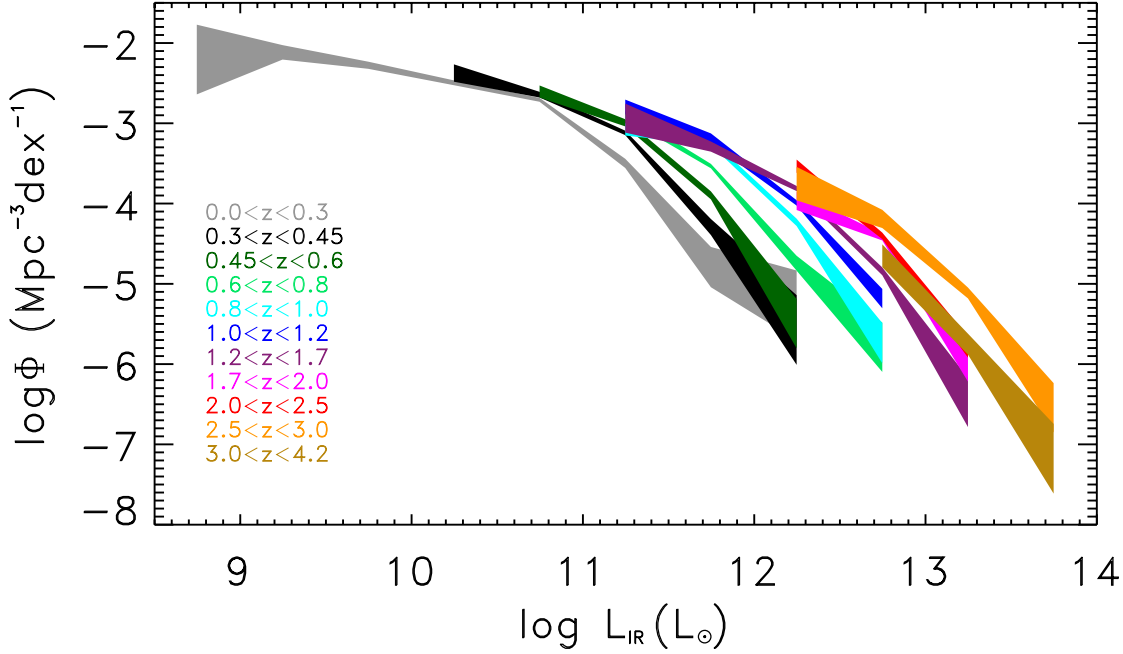


Figure 8. Total IR Luminosity Function estimated by combining the data from the four PEP fields using the Avni & Bahcall (1980) method plotted in all the different redshift intervals considered in this study, from $z \sim 0$ to $z \sim 4$. The different colour-filled areas represent the $\pm 1\sigma$ (Poissonian) uncertainty regions at different redshifts.

tion favours the detection of higher redshift sources than the PEP one, it is more likely that the PEP LF in the higher- z bin is affected by some flux/redshift incompleteness rather than by the presence of low- z sources erroneously placed at high- z by incorrect photometric redshift assignment (catastrophic failures). The values of our total IR LF for each redshift and luminosity bin are reported in Table 6.

3.4 Evolution

In order to study the evolution of the total IR LF, we derive a parametric estimate of the luminosity function at different redshifts. For the shape of the LF we assume a modified-Schechter function (i.e. Saunders et al. 1990), where $\Phi(L)$ is given by

$$\Phi(L) d\log L = \Phi^* \left(\frac{L}{L^*} \right)^{1-\alpha} \exp \left[-\frac{1}{2\sigma^2} \log_{10}^2 \left(1 + \frac{L}{L^*} \right) \right] d\log L, \quad (3)$$

behaving as a power law for $L \ll L^*$ and as a Gaussian in $\log L$ for $L \gg L^*$. The adopted LF parametric shape depends on 4 parameters (α , σ , L^* and Φ^*), whose best fitting values and uncertainties have been found using a non-linear least squares fitting procedure. In detail, while in the first z -bin all the parameters have been estimated, starting from the second z -bin, the values of α and σ have been frozen at the values found at lower redshift, leaving only L^* and Φ^* free to vary (see Table 7). Note that, in the highest redshift bin ($3.0 < z < 4.2$) we are not able to constrain the LF break, while we are up to ~ 3 . Therefore, the results found at this redshift are affected by larger uncertainties than the $z \lesssim 3$ ones. However, although there is some degeneracy in the values of L^* and Φ^* at $3.0 < z < 4.2$, the range of allowed value combinations still giving a reasonable fit to the three

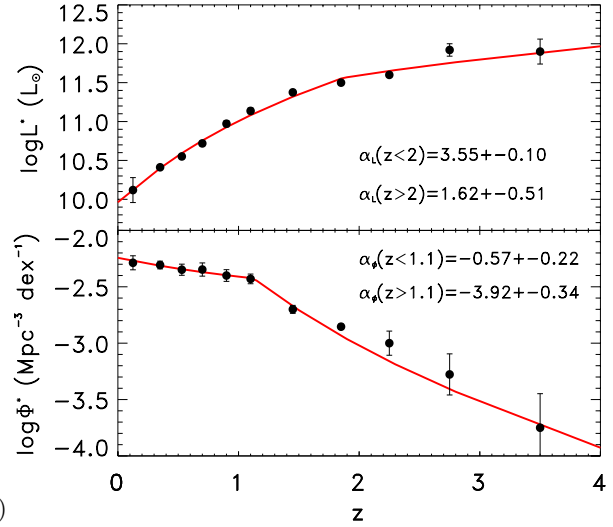


Figure 9. Evolution of L^* and Φ^* as function of z ($L^* \propto (1+z)^{3.55 \pm 0.10}$ at $z \lesssim 1.85$, $\propto (1+z)^{1.62 \pm 0.51}$ at $z \gtrsim 1.85$; $\Phi^* \propto (1+z)^{-0.57 \pm 0.22}$ at $z \lesssim 1.1$, $\propto (1+z)^{-3.92 \pm 0.34}$ at $z \gtrsim 1.1$).

observed data points (constraining the bright-end of the LF) is limited and do not significantly affect our results.

In Fig. 8 we plot the total IR LFs at all redshifts with the $\pm 1\sigma$ Poissonian uncertainty regions (different colours for different z -intervals). There is a clear luminosity evolution with redshift, at least up to $z \sim 3$. The apparent “fall” of the $z > 3$ LF with respect to those at lower redshifts, if real, might imply a global negative evolution of the IR galaxies and/or AGN starting at $z \gtrsim 3$. However, as mentioned above, we must point out that in the highest redshift bin

Table 7. Parameter values describing the curve fitted to the total IR LF

redshift range	α	σ	$\log_{10}(L^*/L_{\odot})$	$\log_{10}(\Phi^*/\text{Mpc}^{-3} \text{ dex}^{-1})$
$0.0 < z < 0.3$	1.15 ± 0.07	0.52 ± 0.05	10.12 ± 0.16	-2.29 ± 0.06
$0.3 < z < 0.45$	1.2^a	0.5^a	10.41 ± 0.03	-2.31 ± 0.03
$0.45 < z < 0.6$	1.2^a	0.5^a	10.55 ± 0.03	-2.35 ± 0.05
$0.6 < z < 0.8$	1.2^a	0.5^a	10.71 ± 0.03	-2.35 ± 0.06
$0.8 < z < 1.0$	1.2^a	0.5^a	10.97 ± 0.04	-2.40 ± 0.05
$1.0 < z < 1.2$	1.2^a	0.5^a	11.13 ± 0.04	-2.43 ± 0.04
$1.2 < z < 1.7$	1.2^a	0.5^a	11.37 ± 0.03	-2.70 ± 0.04
$1.7 < z < 2.0$	1.2^a	0.5^a	11.50 ± 0.03	-3.00 ± 0.03
$2.0 < z < 2.5$	1.2^a	0.5^a	11.60 ± 0.03	-3.01 ± 0.11
$2.5 < z < 3.0$	1.2^a	0.5^a	11.92 ± 0.08	-3.27 ± 0.18
$3.0 < z < 4.2$	1.2^a	0.5^a	11.90 ± 0.16	-3.74 ± 0.30

^a fixed value

PACS data could be affected by incompleteness, since with increasing redshift (and for intrinsically “cold” sources) the true PACS fluxes might fall below the detection limit (i.e. faint sources should be missed even if completeness corrections are perfect), while, if luminous enough, they can still be detectable by SPIRE. This effect is expected to be more relevant in COSMOS, where SPIRE data are quite deep relatively to PACS, while it should not happen in GOODS-S, where PACS data are very deep compared to the SPIRE ones (which are limited by confusion). The high- z incompleteness of PACS surveys might also be emphasised by the redshift incompleteness of the COSMOS sample, that could affect the highest redshift bins more than the lower ones (although the redshift incompleteness seems to be independent of redshift; see Berta et al. 2011). However, we must point out that a decrease at $z \gtrsim 2.7-3$ similar to that observed in our data, is also observed in the space density of X-ray (see Brusa et al. 2009; Civano et al. 2011) and optically selected AGN (Richards et al. 2006), and of sub-mm galaxies (Wall et al. 2008), as well as in the HerMES total IR LF – though less evident – (from 250- μm data; Vaccari et al., in preparation; see Fig. 7). Moreover, our result is in agreement with the recent finding of Smit et al. (2012), that the characteristic value of the galaxy SFR exhibits a substantial, linear decrease as a function of redshift from $z \sim 2$ to $z \sim 8$.

In Fig. 9 we show the values of L^* and Φ^* at different redshifts, with the best curve ($\propto (1+z)^k$) fit to the data points. The values of the curve slopes and of the redshifts corresponding to evolutionary breaks have been chosen to be those which minimise the χ^2 of the fit with two power-laws. We find a significant variation of L^* with z , which increases as $(1+z)^{3.55 \pm 0.10}$ up to $z \sim 1.85$, and as $(1+z)^{1.62 \pm 0.51}$ between $z \sim 1.85$ and $z \sim 4$. The variation of Φ^* with redshift starts with a slow decrease as $(1+z)^{-0.57 \pm 0.22}$ up to $z \sim 1.1$, followed by a rapid decrease $\propto (1+z)^{-3.92 \pm 0.34}$ at $z \gtrsim 1.1$ and up to $z \sim 4$.

Previous estimations of the evolution of L^* and Φ^* (i.e. Caputi et al. 2007, Bethermin et al. 2011 and Marsden et al. 2011) discussed a decrease in the density of far-IR sources between $z=1$ and $z=2$. In particular, Bethermin et al. (2011) and Marsden et al. (2011), by evolving a parameterised far-IR LF, explored the evolution required by the source counts in the parameter space. The results of these works (especially those of Bethermin et al. 2011, showing a decrease

of Φ^* at $z > 1$ and a flatter trend on the evolution of L^* at $z > 2$), are close to ours, though with the source counts only it was not possible to constrain the evolution of IR sources at $z > 2$.

3.5 Evolution of the Different IR Populations

The PEP survey, given its size and its coverage in luminosity and redshift, allows us to go further in investigating the evolution of the total IR LF: it gives us the opportunity to study the evolution of the different galaxy classes that compose the global IR population. To investigate the different evolutionary paths of the various IR populations, we have computed the $1/V_{\text{max}}$ LFs separately for the five galaxy classes defined by the SED-fitting analysis. In Fig. 10 we show the total IR LFs derived from the combined PEP samples for the different SED classes (coloured filled areas). The results of the fit to a modified Schechter function (see equation 3) for each population are overplotted on the data. In fact, by following the same procedure adopted for the global luminosity function, a parametric fit to the LFs at different redshifts has been performed also for the single populations. The α and σ parameters, for each population, have been estimated at the redshift where the corresponding LF is best sampled (not necessarily at the lowest z -bin as for the global LF). Subsequently, the values of α and σ have been frozen at the values found in the “optimal” redshift bin, leaving only L^* and Φ^* free to vary. In Fig. 11, analogously to Fig. 9, we show the values of L^* and Φ^* at different redshifts for the different populations, with the best least square fitting curves ($\propto (1+z)^k$) overplotted. The best-fitting values of α , σ , L^* and Φ^* in the first redshift bin, together with the parameters describing the luminosity ($k_{L,1}$, $k_{L,2}$ and $z_{b,L}$: $\propto (1+z)^{k_{L,1}}$ to $z=z_{b,L}$, $\propto (1+z)^{k_{L,2}}$ at $z > z_{b,L}$) and density evolution ($k_{\rho,1}$, $k_{\rho,2}$ and $z_{b,\rho}$: $\propto (1+z)^{k_{\rho,1}}$ out to $z=z_{b,\rho}$, $\propto (1+z)^{k_{\rho,2}}$ at $z > z_{b,\rho}$) are reported in Table 8.

A clear result of our analysis is that the evolution derived for the global IR LF is indeed a combination of different evolutionary paths: the far-IR population does not evolve all together “as a whole”, as it is often assumed in the literature, but is composed by different galaxy classes evolving differently and independently. As shown in Fig. 10, the normal **spiral** galaxy population dominates the luminosity function at low- z , from the local Universe up to $z \sim 0.5$. Mov-

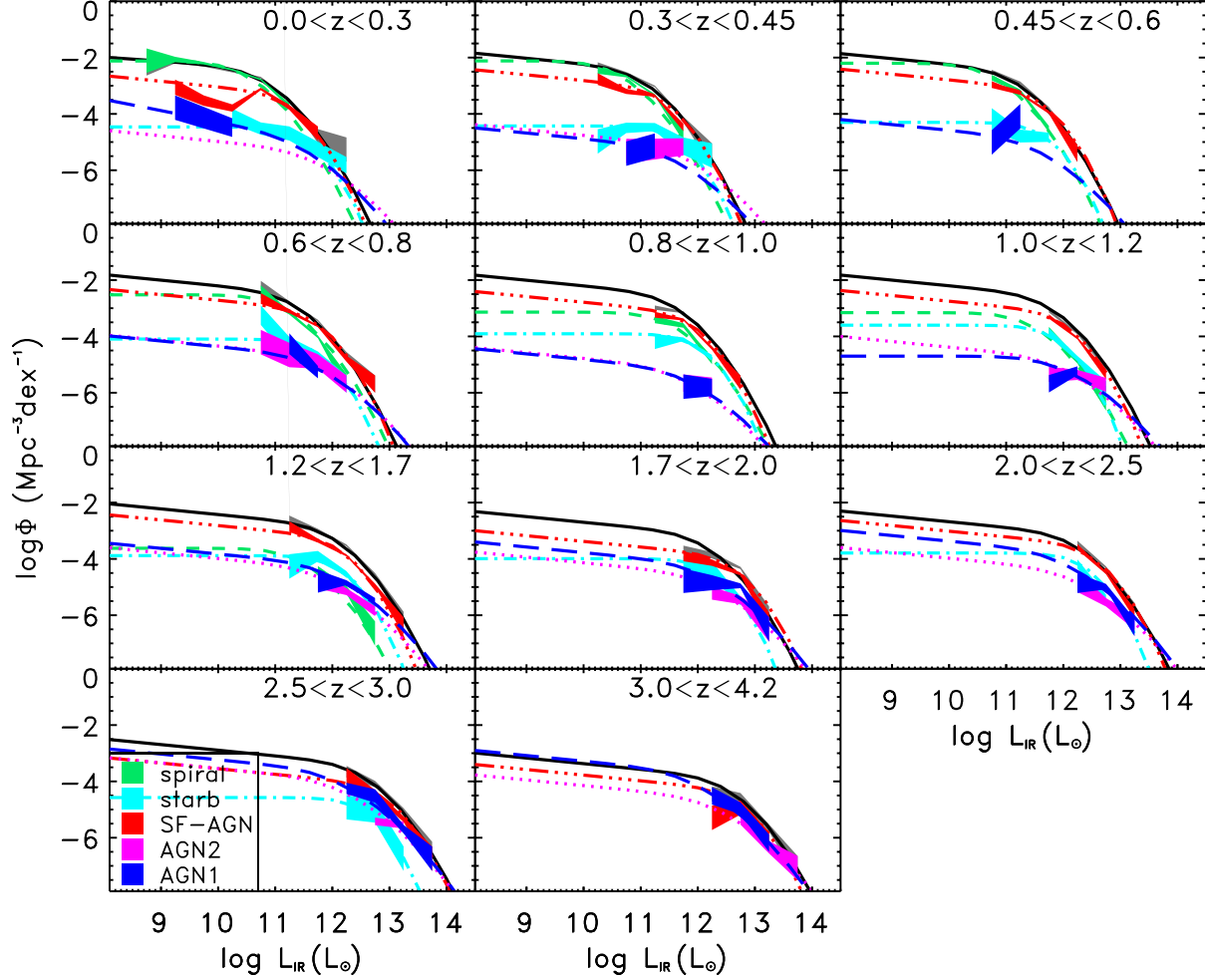


Figure 10. Total IR Luminosity Function estimated with the $1/V_{\max}$ method by combining the data from the four PEP fields for the different populations (their $\pm 1\sigma$ uncertainty regions are shown as coloured filled areas: green for **spirals**; cyan for **starbursts**; red for **SF-AGN**; magenta for **AGN2**; and blue for **AGN1**), compared to the total LF ($\pm 1\sigma$, grey filled area, same as in Fig. 7). The best-fit modified Schechter functions are also plotted, extrapolated to fainter and brighter luminosities than covered by the data (with the black curve being for the total LF and the same colours used for the filled areas as for the single populations).

Table 8. Parameter values describing the curve fitted to the total IR LF of the different SED populations

	α	σ	$\log_{10}(L^*/L_{\odot})$	$\log_{10}(\Phi^*/\text{Mpc}^{-3} \text{ dex}^{-1})$	$k_{L,1}$	$k_{L,2}$	$z_{b,L}$	$k_{\rho,1}$	$k_{\rho,2}$	$z_{b,\rho}$
	$(0.0 < z < 0.3)$									
spiral	1.00 ± 0.05	0.50 ± 0.01	9.78 ± 0.04	-2.12 ± 0.01	4.49 ± 0.15	0.00 ± 0.46	1.1	-0.54 ± 0.12	-7.13 ± 0.24	0.53
starburst	1.00 ± 0.20	0.35 ± 0.10	11.17 ± 0.16	-4.46 ± 0.06	1.96 ± 0.13			3.79 ± 0.21	-1.06 ± 0.05	1.1
SF-AGN	1.20 ± 0.02	0.40 ± 0.10	10.80 ± 0.02	-3.20 ± 0.01	3.17 ± 0.04			0.67 ± 0.05	-3.17 ± 0.15	1.1
AGN2	1.20 ± 0.20	0.70 ± 0.20	10.80 ± 0.20	-5.14 ± 0.17	1.41 ± 0.33			2.65 ± 0.32		
AGN1	1.40 ± 0.30	0.70 ± 0.20	10.50 ± 0.20	-5.21 ± 0.11	1.31 ± 0.02			3.00 ± 0.25		

ing to higher redshifts, the number density of galaxies with spiral galaxy SEDs sharply decreases, while their luminosity continues to increase, at least up to $z \sim 1$ (see the Φ^* and L^* parameter trends shown in Fig. 11). We note that what we observe between $z \sim 0$ and $z \sim 1$ for the **spiral** SED galaxies is an increase of L^* by a factor of ~ 5 , and a decrease of Φ^* by a factor of ~ 10 . Since the two evolutions are not independent, the “total” evolutionary effect results from the

combination of the two (as can be observed in the total IR luminosity density, see Section 4). A way to derive the “total” effect of evolution on a LF is to fix at a given volume density value and see how the luminosity corresponding to that value changes: indeed we find an increase by a factor of ~ 2.5 between $z=0$ and $z=1$ for the **spiral** LF, in good agreement with previous results, either empirical (for morphologically classified disk galaxies; Scarlata et al. 2007) or

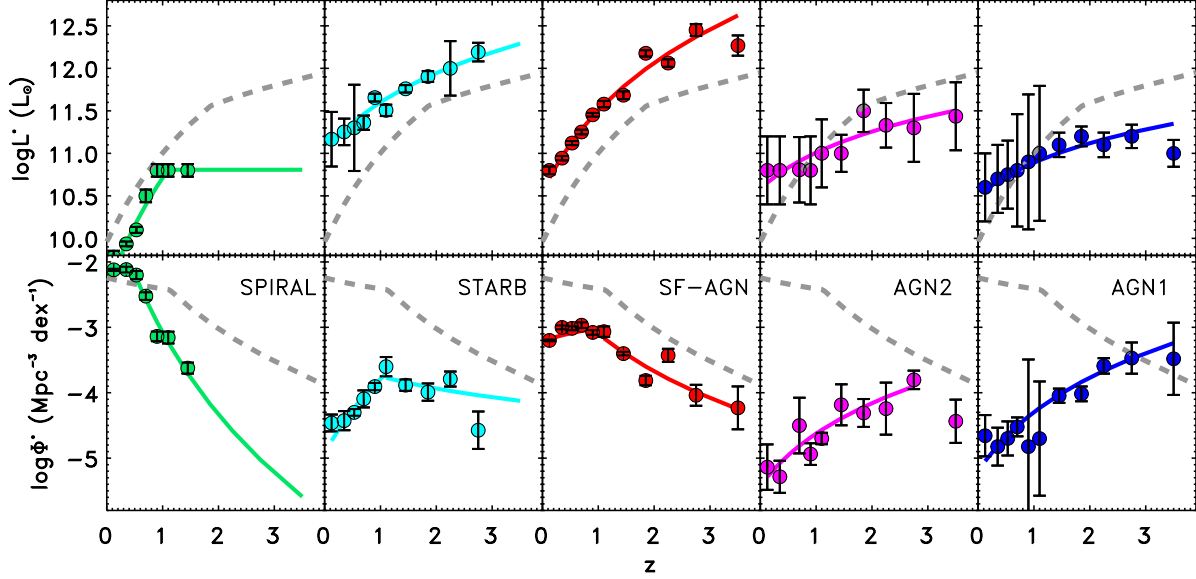


Figure 11. Evolution of L^* (top) and Φ^* (bottom) as a function of redshift (in the form $\propto (1+z)^\kappa$) for the different IR populations. For comparison, the L^* and Φ^* evolution of the “global” total IR LF (plotted in Fig. 9), are shown as grey dashed lines.

theoretical (from chemical evolution models of Milky Way-like galaxies; Colavitti, Matteucci & Murante 2008). Over the whole redshift range 0.5–3, the “total” luminosity function is dominated by the SF-AGN population. The number density of SF-AGN is nearly constant from the local Universe up to $z \sim 1.5$, showing a slight decrease at higher redshifts, while their luminosities show positive evolution up to the highest redshifts ($z \sim 3.5$ –4). From Fig. 10 we note a sort of bimodality in the SF-AGN LFs (at $z \lesssim 0.45$, where we are able to cover a larger luminosity range). This bimodality is indeed to be ascribed to the crossing of two contributions: that of the SF-AGN(Spiral) population, responsible for the faint-end steepness of the LFs, and that of the SF-AGN(SB) population, dominating the bright-end of the SF-AGN LFs and declining at low L_{IR} (not reported in the figure).

The starburst galaxy population never dominates. The redshift range where we observe the highest contribution from the starburst galaxies is at $z \sim 1$ –2, while in the local Universe their contribution is almost negligible (i.e. their Φ^* parameter shows an opposite trend with respect to that of spiral galaxies see Fig. 11).

The AGN1 and AGN2 populations show a very similar evolutionary trend as a function of z , both in Φ^* and L^* . These powerful AGN populations dominate only the very bright end of the total IR LF, although their number densities and luminosities keep increasing from the local Universe up to the higher redshifts. At $z > 2.5$ the AGN1 and AGN2 populations become as important as the SF-AGN one, with the total IR LF of PACS-selected sources in the redshift range 2.5–4 being totally dominated by objects containing an AGN.

3.6 Total IR LF in Mass and Specific Star-Formation Rate bins

3.6.1 Stellar masses and SFR from SED fitting

The wealth of multi-wavelength data available in the cosmological fields included in our work allow us to perform a detailed SED fitting of all sources, in order to derive their most relevant physical parameters (e.g. stellar masses). To derive stellar masses we have fitted the broad-band SEDs of our sources using a modified version of MAGPHYS (Da Cunha et al. 2008), which is a code describing the SEDs using a combination of stellar light and emission from dust heated by stellar populations. In particular, the MAGPHYS software simultaneously fits the broad-band UV-to-far IR observed SED of each object, ensuring an energy balance between the absorbed UV light and that re-emitted in the far-IR regime. The main assumptions are that the energy re-radiated by dust is equal to that absorbed, and that starlight is the only significant source of dust heating. We refer to Da Cunha et al. (2008) for a thorough formal description of how galaxy SEDs are build. At each source’s redshift, the code chooses among different combinations of star formation histories, metallicities and dust contents, associating a wide range of optical models to a wide range of infrared spectra and comparing to observed photometry, seeking for χ^2 minimization. Each star formation (SF) history is parameterised in terms of an underlying continuous model with exponentially declining star formation rate (SFR), on top of which are superimposed random bursts (see Da Cunha et al. 2008, Da Cunha et al. 2010). We note that, although the MAGPHYS assumption of exponentially declining SFR might not be the best to reproduce the SFR history of $z > 1.5$ star-forming galaxies (i.e. exponentially increasing or increasing SFR would be better choices, as widely discussed by Maraston et al. 2010 and Reddy et al. 2012), in our specific case it does not affect the results. In fact, we do not use the MAGPHYS derived SFRs, but we compute them by integrating the best-fitting

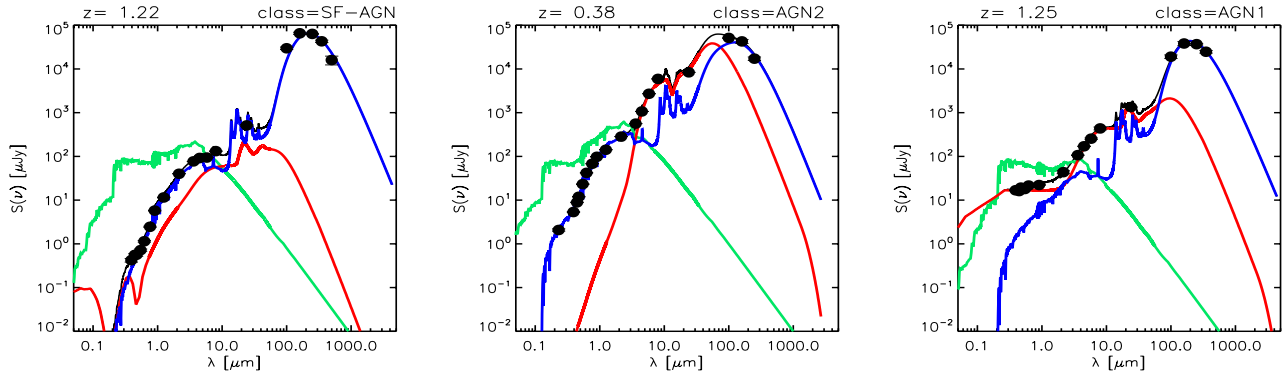


Figure 12. Examples of PEP AGN SEDs fitted by the MAGPHYS code (Da Cunha et al. 2008) modified by Berta et al. (2013) to add an AGN component from the Fritz et al. (2006) library. From *left to right* the result of the fit to a SF-AGN, an AGN2 and an AGN1 are shown. The stellar component unattenuated by dust is shown in green, while the dust-attenuated spectrum with dust IR re-emission is shown in blue. The AGN contribution is shown in red, while the black curve is the total fitted spectrum, obtained from the sum of the blue and red components. Our data are represented by the black dots.

SED (resulting from *Le Phare*). The models are distributed uniformly in metallicity between 0.2 and 2 times solar. Since the MAGPHYS code assumes that starlight is the only significant source of dust heating, thus ignoring the presence of a possible AGN component, Berta et al. (2013) have developed a modified version of the MAGPHYS code by adding a torus component to the modelled SED emission, combining the Da Cunha et al. (2008) original code with the Fritz et al. (2006) AGN torus library (see also Feltre et al. 2012). The spectral fitting is performed by comparing the observed SED of our galaxies to every model in the generated library, at the corresponding redshift. A χ^2 minimisation provides the quality of each fit. We must point out that the mass derivation for unobscured AGN (i.e. AGN1) is a problematic issue, therefore the masses estimated for that class of objects are the most uncertain ones. One source of uncertainty in the mass measurement for AGN1 is due to the fact that, in these objects, the UV part of the SED is likely dominated by the AGN rather than by the host galaxy. This may produce an underestimate of the mass, since, if the AGN contribution is not taken into account, the data can be fit by a bluer, younger and smaller mass object. On the other hand, the mid-IR part of the AGN SED is dominated by dust emission from the dusty torus heated by the central black hole. If a proper decomposition into a stellar and a torus component is not performed, the use of a pure stellar template for estimating the mass from SED-fitting tends to reproduce the mid-IR emission with an older, redder and larger mass object (mass sometimes larger by a factor of 2 than those derived through a decomposition procedure; Santini et al. 2012). These two effects lead to an increase in the uncertainty of the mass derivation, although they might somewhat compensate their effects for a large sample of objects. For this reason, we obtained measurements of the stellar masses of our objects containing an AGN by means of the specific decomposition technique developed by Berta et al. (2013), to separate stellar and nuclear emission components. Examples of the results of this decomposition applied to SF-AGN, AGN2 and AGN1 are shown in Fig. 12. Masses of AGN estimated with the original MAGPHYS and with the Berta et al. (2013) code have been compared, showing very good agreement and small dispersion around the 1–1 relation. Similarly,

as further check, we have also computed stellar masses with different code (*Hyperz*; Bolzonella et al. 2000) and stellar library (BC03, Bruzual & Charlot 2003, instead of the CB07 used as default by MAGPHYS), finding good agreement and no systematics, too.

We have derived the SFRs from the total IR luminosities (estimated from the SED fitting described in Section 2.1) with the standard Kennicutt (1998) relation (converted to Chabrier IMF), after subtracting the AGN contribution to L_{IR} . We note that the total IR luminosity in PEP sources is usually dominated by star formation, even in objects for which an AGN dominates the optical/near-IR/mid-IR part of the spectrum.

3.6.2 LFs in different mass bins

We compute the total IR LF for galaxies of different stellar masses: $8.5 \leq \log(M/M_\odot) < 10$, $10 \leq \log(M/M_\odot) < 11$, and $11 \leq \log(M/M_\odot) < 12$, by means of the standard $1/V_{\text{max}}$ formalism, and we show the results in Figs. 13 (total IR LF in different z -bins) and 14 (ratio between the LF in the mass intervals and the total IR LF).

We have compared our results with the SFR function (SFR converted to IR luminosity using the Kennicutt (1998) relation) of massive ($\log(M/M_\odot) > 10$) galaxies derived by Fontanot et al. (2012) from the GOODS-MUSIC sample at redshift $0.4 < z < 1.8$. In the common redshift and luminosity range we find an excellent agreement with our total IR LF, which is dominated by sources with $10 < \log(M/M_\odot) < 11$. At lower luminosities, not sampled by our data, the Fontanot et al. (2012) LFs are characterized by a double-peaked structure, interpreted in terms of the well-known bimodality in the colour(SFR)-Mass diagram. As expected from the SFR-Mass relation, the knee (L^*) of our IR LFs in different mass bins moves to higher luminosities with increasing masses (i.e. at $0.0 < z < 0.3$, $\log(L^*/L_\odot)$ changes from ~ 9.5 for $\log(M/M_\odot) = 8.5-10$ sources, to ~ 11.3 for sources with $\log(M/M_\odot) > 11$).

The slope of the total IR LF in each mass bin is always similar to (or flatter than) the “global” LF (total, including all the masses). The lower mass galaxies dominate at lower luminosities ($\log(L/L_\odot) < 9$), while the most massive

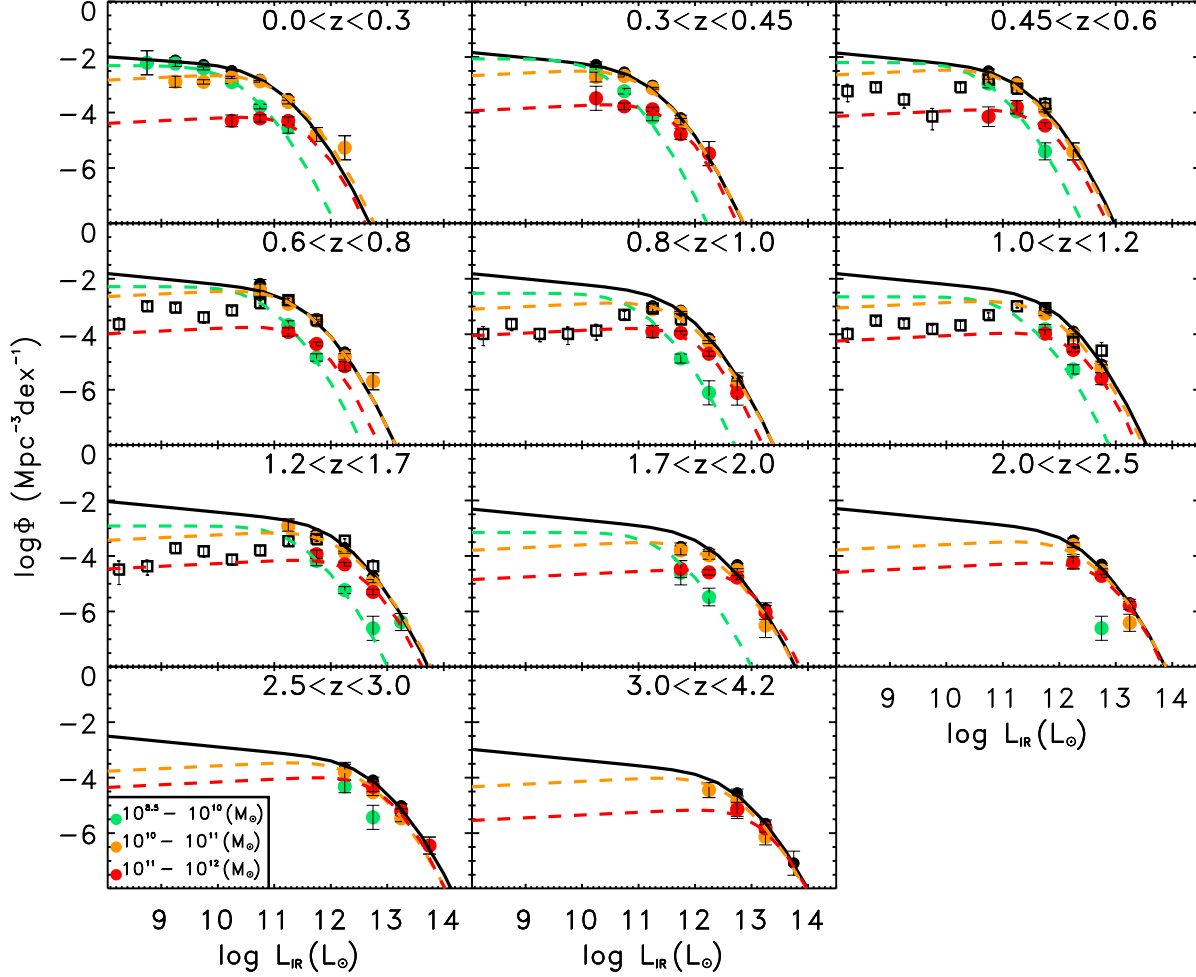


Figure 13. Contribution to the total IR LF from galaxies in three different mass intervals: $8.5 < \log(M/M_{\odot}) < 10$ (green); $10 < \log(M/M_{\odot}) < 11$ (orange); and $11 < \log(M/M_{\odot}) < 12$ (red). For comparison, the SFR function of $\log(M/M_{\odot}) > 10$ sources in the GOODS-S by Fontanot et al. 2012 (converted to total IR LF) is shown as black open squares.

galaxies ($\log(M/M_{\odot}) > 11$) contribute only at higher luminosities, even if they never dominate the LF. At all masses, the LF evolves with redshift, following the evolution of the “global” LF. Fig. 14 shows that the main contribution (>50 per cent) to the total IR LF is due to intermediate-mass objects ($\log(M/M_{\odot}) = 10$ –11) at all redshifts and luminosities, with their fraction remaining almost the same from $z = 0$ to $z \sim 4$, simply shifting to higher luminosities. Lower mass objects ($\log(M/M_{\odot}) = 8.5$ –10) contribute significantly only at $\log(L/L_{\odot}) < 10$, with their fraction just shifting to higher luminosities with redshift, but always being below 30 per cent at $z > 0.45$ and $\log(L/L_{\odot}) > 11$. The contribution of the most massive objects ($\log(M/M_{\odot}) = 11$ –12) increases with IR luminosity and redshift, becoming significant (>50 per cent) only at $z > 1.7$ and $\log(L/L_{\odot}) > 12.5$. Thus the bulk of the IR luminosity is produced by star-forming galaxies of mass around the characteristic mass M^* of the Schechter mass function.

3.7 Specific-SFR and the main sequence of star-forming galaxies

Having computed stellar masses and SFRs for each source, we can check how the PACS selected sources populate the SFR–stellar mass plane and the so called main sequence (MS) of star-forming galaxies, as a function of redshift. This relation (i.e. SFR versus stellar mass) has been shown to be quite tight in the local Universe (Peng et al. 2010, 2011) and well established at redshift $z \sim 1$ (Elbaz et al. 2007) and up to $z \sim 2$ (Daddi et al. 2007, Rodighiero et al. 2011) and $z \sim 3$ (Magdis et al. 2010), with normalisation scaling as $\sim (1+z)^{2.8}$ out to $z \sim 2$, as shown by Sargent et al. (2012) (see also Elbaz et al. 2007, Rodighiero et al. 2010b, Pannella et al. 2009, Karim et al. 2011). At $z \sim 2$ and ~ 1.5 we assume a slope of 0.79 for the MS in the SFR versus stellar-mass plane (according to Rodighiero et al. 2011 and Sargent et al. 2012), while at $z \sim 1$ we assume a slope of 0.9, as found by Elbaz et al. (2007), and we limit our investigation to the redshift range $0.8 < z < 2.2$.

By combining UV and far-IR data, Rodighiero et al. (2011) re-evaluated the locus of the MS at $z \sim 2$, showing that objects lying a factor of 4 above the MS (in SFR) can be con-

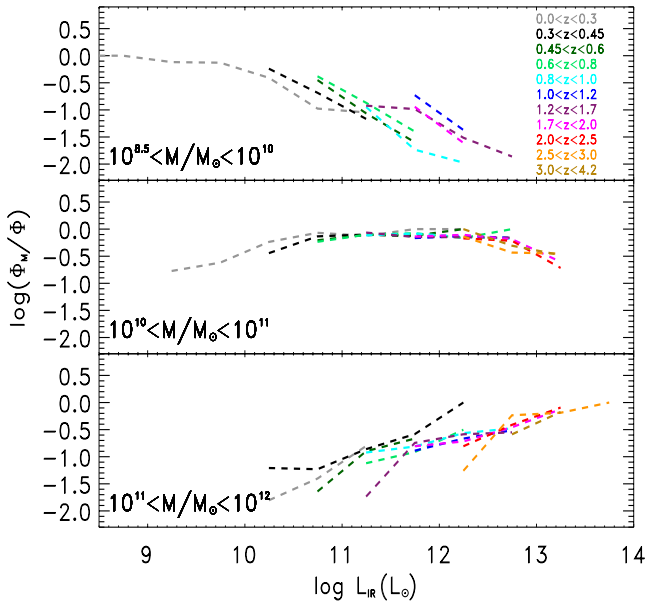


Figure 14. Ratio between the IR LF of galaxies in three different mass intervals: $8.5 < \log(M/M_\odot) < 10$ (top); $10 < \log(M/M_\odot) < 11$ (middle); and $11 < \log(M/M_\odot) < 12$ (bottom), as shown in Fig. 13, and the total IR LF, plotted in all the different redshift intervals considered in this study, from $z \sim 0$ to $z \sim 4$. The different colours represent the different redshifts, as explained in the plot.

considered as outliers with respect to the average locus where smoothly star-forming galaxies spend most of their lives in a secular and steady regime. In that work, off-sequence sources (characterized by very high specific-SFRs) are assumed to be in a starburst mode, and are found to contribute only 2 per cent of mass-selected star-forming galaxies and to account for only 10 per cent of the cosmic SFR density at $z \sim 2$ (Rodighiero et al. 2011).

In order to check what kind of objects we could classify as on- and off-MS sources in our IR sample compared to previous findings, based either on IR or on optical surveys (e.g. Rodighiero et al. 2011; Sargent et al. 2012), we have splitted our sample into off-MS and on-MS. For consistency with previous studies we have applied the same criterion as Rodighiero et al. (2011) (0.6 dex above the MS) over the whole $0.8 < z < 2.2$ redshift range, by using as a reference MS the relation found by Rodighiero et al. (2011) at $z \sim 2$, scaled as described above at $z \sim 1.5$, and the relation found by Elbaz et al. (2007) at $z \sim 1$. In Fig. 15, we show the SFR versus stellar mass distributions in three redshift bins ($0.8 < z < 1.25$, $1.25 < z < 1.8$ and $1.8 < z < 2.2$), for the PACS sources included in the computation of the luminosity functions presented in this work. The colour code marks the different SED-classes to which each source belongs. We also report the typical loci of the MS at the various redshifts (scaling as $(1+z)^{2.8}$, as mentioned above). Details are given in the caption of the Figure. The typical far-IR selection bias (PACS-*Herschel* in this case) appears as an approximate horizontal SFR cut (Rodighiero et al. 2011, Wuyts et al. 2011), shown as thin dotted line in Fig. 15. We note that the trends of mid/far-IR SEDs with offset from the main sequence observed in Fig. 15 (and widely discussed in Section 5) are in good agreement

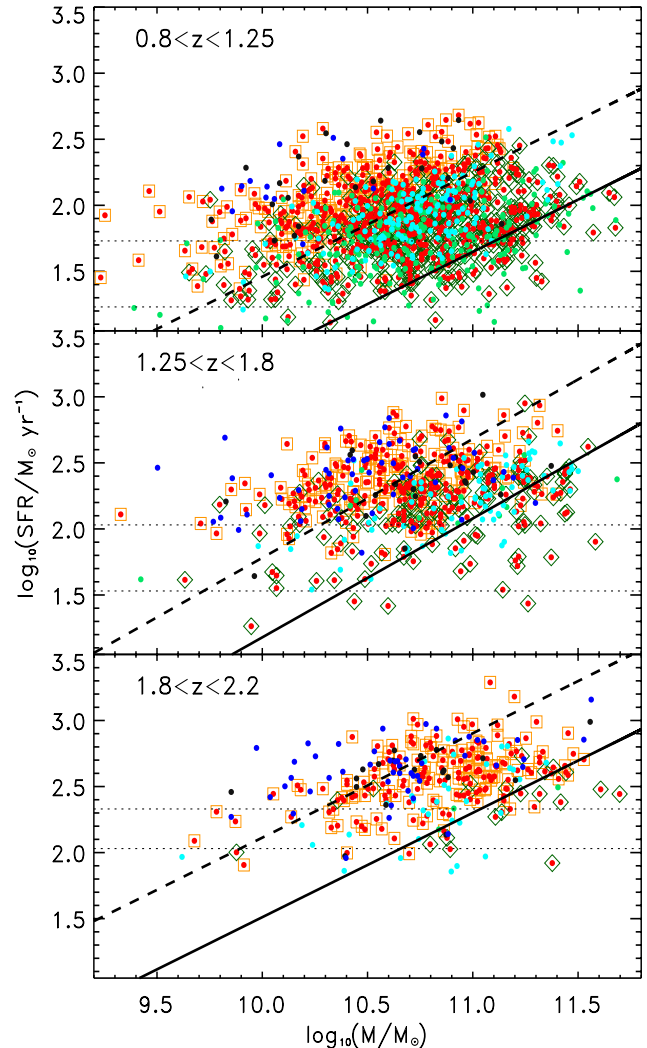


Figure 15. SFR versus stellar mass for our PEP 160- μm sources (green, **spiral**; cyan, **starburst**; red, **SF-AGN**; magenta, **AGN2**; blue, **AGN1**), in three redshift bins (from left to right): $0.8 < z < 1.25$; $1.25 < z < 1.8$; and $1.8 < z < 2.2$. The relation known as main sequence is plotted as a solid line (from Elbaz et al. 2007 in the lower redshift bin, rescaled as $(1+z)^{2.8}$ in the central bin and from Rodighiero et al. 2011 at $z \sim 2$), while the dashed-line shows the same relation 0.6 dex higher, indicating the separation between MS and above MS objects adopted by Rodighiero et al. (2011). The horizontal dotted lines show the nominal SFR limits of the GOODS-S (lower) and COSMOS (upper) samples in the different redshift intervals. The orange open squares and the dark-green open diamonds mark the two sub-classes of SF-AGN galaxies: the SF-AGN(SB) and SF-AGN(Spiral) respectively.

with the results of Elbaz et al. (2011) and Nordon et al. (2012).

With the selection based on Rodighiero et al. (2011) and overplotted in Fig. 15 (sources qualify as “off-MS” if they lie more than 0.6 dex above the observed SFR-stellar mass relation), we can compute the contribution of off-MS (also called “starburst” in the literature) and on-MS (“steady star-formers” in the literature) sources to the total IR LFs. This is presented in Fig. 16, where the total IR LFs of on- and off-MS sources have been computed indepen-

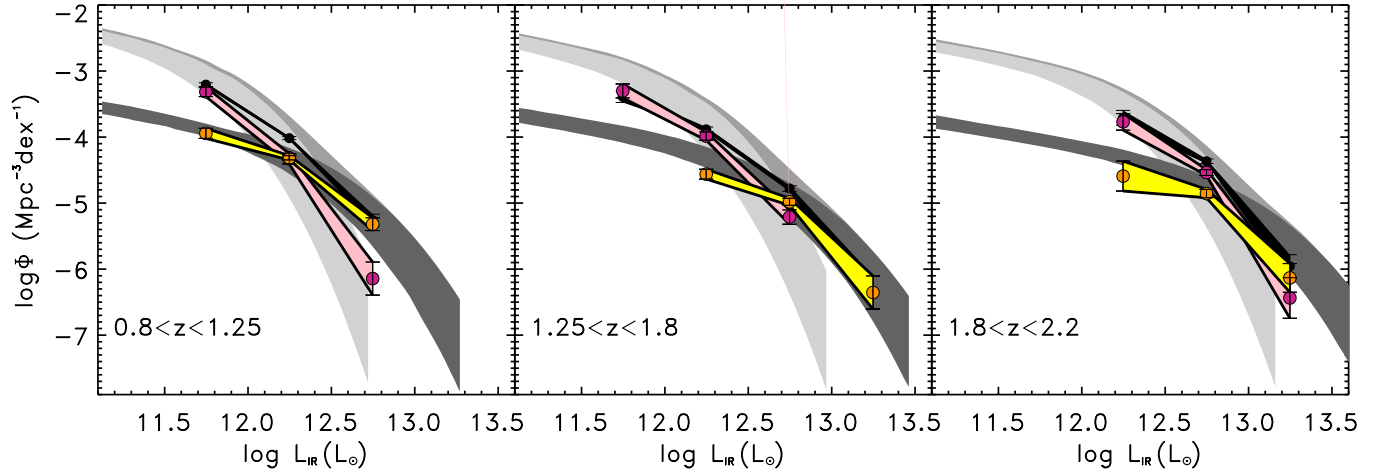


Figure 16. Contribution of MS ($\pm 1\sigma$ uncertainty region: pink shaded area) and off-MS ($\pm 1\sigma$ uncertainty region: yellow shaded area) galaxies to the total IR LF (black filled dots and area) in three different redshift bins. For comparison, the recent estimates of Sargent et al. (2012) in similar z -bins are shown as grey (total LF, in background), light grey (MS LF) and dark grey (off-MS LF) filled regions.

dently for the three redshift bins. The pink and yellow filled areas correspond to the $\pm 1\sigma$ uncertainty regions of the total IR LFs estimated for the on-MS and off-MS populations, respectively, while the black filled area marks the global population. Our results are compared with the recent estimates by Sargent et al. (2012), a non-parametric approach that is based on three basic observables: the redshift evolution of the stellar mass function for star-forming galaxies; the evolution of the sSFR of MS galaxies; and a double-Gaussian decomposition of the sSFR distribution at fixed stellar mass into a contribution (assumed redshift- and mass-invariant) from MS and off-MS (i.e. starburst) activity. The evolution of the two populations found both in our data and the Sargent et al. (2012) model are very similar. Both data and estimates indicate that the bright-end of the total IR LF is dominated by off-MS sources. However, although consistent within the uncertainties, the relative contribution of off-MS sources seems to be stronger in the Sargent et al. (2012) model than observed in the present computation, where we find that the bright-ends of the PEP IR LFs are more similarly populated by MS and off-MS sources (especially at $z \sim 2$). This difference can be at least partly ascribed to the sharp cut we apply to separate MS from off-MS sources, while Sargent et al. (2012) model the off- and on-MS populations with two continuous log-normal distributions centred at 0.6 dex above the MS and exactly on the MS respectively, of which the one describing the “starburst” population has wings that extend into our on-MS selection region (hence attributing more sources to the starburst category than are selected in our off-MS class). A better agreement between our data and the estimates of Sargent et al. (2012) is found at the faint-end of the LFs that appear to be completely dominated by the “normal” MS galaxies at all redshifts (although the total and relative contributions at $\log(L_{\text{IR}}/L_{\odot}) < 12$ are slightly lower in the data than predicted by Sargent et al. (2012). Good agreement is also found with respect to the evolution of the cross-over luminosity (i.e. where the contributions from on- and off-MS sources are equal); in the model of Sargent et al. (2012) the cross-over luminosity simply shifts to higher luminosities and lower densities (according to the

luminosity and density evolution considered). Note that the model assumptions rely on results from different surveys, selected at different wavelengths, complete in mass and with good sampling of the MS. On the other hand, our selection is in SFR and our sources do not follow any clear sequence in stellar mass–SFR plane, because, except at the highest masses, the data are not deep enough to reach well into the main sequence. These different selection effects are likely to lead to some differences between our LFs and the estimates of Sargent et al. (2012).

To quantify the relative contribution of the two populations, our observed data have been fitted with a modified Schechter function, in order to integrate them and compute their comoving number and luminosity densities as functions of redshift (see next Section).

4 NUMBER DENSITY AND IR LUMINOSITY DENSITY

We derive the evolution of the comoving number and luminosity density (total, see Fig 18) of the PEP sources, either belonging to the different SED classes (Fig 18, *left*), to the on- and off-MS categories (*middle*) and to the different mass intervals (*right*), by integrating the total IR LF in the different redshift bins from $z \sim 0$ to $z \sim 4$. To compute the number (and IR luminosity) density, we integrate the Schechter functions that best reproduce the different populations/mass/sSFR-classes, down to $\log(L/L_{\odot}) = 8$. We note that here we consider lower limits the number and luminosity densities at $3.0 < z < 4.2$, since our LF estimate in that redshift bin is likely to be incomplete, as discussed in Section 3.3. We find that the number density of the whole IR population is nearly constant in the $z = 0$ – 1.2 redshift range (slightly increasing from $z \sim 0$ to $z \sim 0.5$), decreasing at $z > 1.2$ (see *top* panels of Fig. 18). When decomposing the number density according to the different SED classes, we observe that normal **spiral** galaxies dominate the local density, with a smaller contribution also from the **SF-AGN** population and a negligible one from **starburst**, **AGN1** and **AGN2**. The space

density of **spiral** galaxies decreases rapidly at $z \gtrsim 0.5$, while that of **SF-AGN** stays nearly constant at $0.5 \lesssim z \lesssim 2.5$, largely dominating in that redshift range. **Starburst** galaxies never dominate, while the number density of the bright AGN (both **AGN1** and **AGN2**) increases with redshift, from $\sim 10^{-4} \text{ Mpc}^{-3}$ at $z \sim 0$ to $\sim 1\text{--}2 \times 10^{-3} \text{ Mpc}^{-3}$ at $z \sim 3$. At higher redshifts the AGN population largely dominates the number density. If the overall contribution to the IR luminosity density (ρ_{IR}) from the AGN components of galaxies is small, ρ_{IR} can be considered as a proxy of the SFR density (ρ_{SFR}). As a further check, we have therefore studied the evolution of the **SF-AGN** population (which dominates the distribution of sources) by dividing this class into **SF-AGN(SB)** and **SF-AGN(Spiral)** sub-classes and studying their evolution separately. Indeed, we have found different evolutionary paths for the two populations, the former dominating at higher redshifts and showing a behaviour similar to that of AGN-dominated sources (e.g. **AGN1** and **AGN2**), the latter dominating at intermediate redshifts (between $z \sim 1$ and 2), rising sharply from $z \sim 2$ toward the lower redshifts and decreasing, while the **spiral** population rises at $z \lesssim 1$. These evolutionary trends, in terms of number and luminosity density, have been reported in Fig. 18 as orange dot-dot-dot-dashed (**SF-AGN(SB)**) and dark-green dashed (**SF-AGN(Spiral)**) curves.

Galaxies following the SFR–mass relation are always dominant over the off-MS population, at all redshifts (although their space density decreases with increasing z , as well as the “global” number density), while the number density of the latter population remains nearly constant between $z \sim 0.8$ and $z \sim 2.2$.

In all the mass bins, the trends with redshift of the galaxy number densities are similar to the “global” one, decreasing at higher redshifts, although with slightly different slopes for the different mass intervals. The number densities of low mass galaxies ($8.5 < \log(M/M_{\odot}) < 10$), reported in the top right panel of Fig. 18, have been computed by integrating the best-fitting modified Schechter function only to $z \sim 2$, since data were not enough to derive reliable fits at higher redshifts. To this redshift, these sources outnumber the higher mass ones, although they fall steeply above $z \sim 1$, when they reach about the same volume density of higher mass galaxies ($10 < \log(M/M_{\odot}) < 11$). Massive objects ($\log(M/M_{\odot}) > 11$) never dominate (always below 5 per cent) the total number density.

The total IR LF allows a direct estimate of the total comoving IR luminosity density (ρ_{IR}) as a function of z , which is a crucial tool for understanding galaxy formation and evolution. Although ρ_{IR} can be converted to a SFR density (ρ_{SFR}) under the assumption that the SFR and L_{IR} quantities are connected by the Kennicutt (1998) relation, before doing that we must be sure that the total IR luminosity is produced uniquely by star-formation, without contamination from an AGN. The SED decomposition and separation into AGN and SF contributions show a negligible contribution to L_{IR} (< 10 per cent) from the AGN in most of the **SF-AGN**, and a SF component dominating the far-IR even in the majority of more powerful AGN (**AGN1** and **AGN2**). Here we prefer to speak in terms of ρ_{IR} rather than of ρ_{SFR} , since, especially at high redshift – where the AGN-dominated sources are more numerous – the conversion of ρ_{IR} could represent only an upper limit to ρ_{SFR} . Note, how-

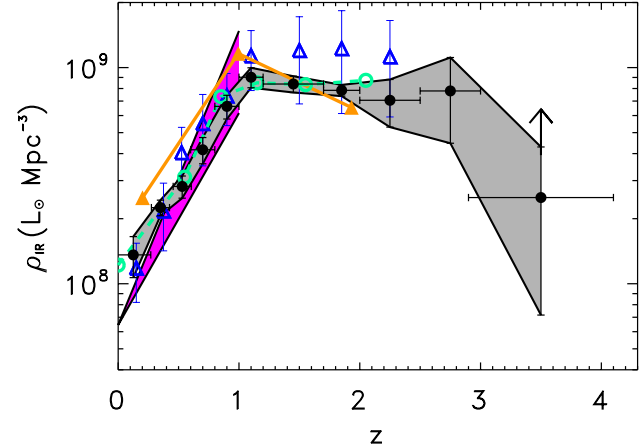


Figure 17. Redshift evolution of the total IR luminosity density (ρ_{IR} , obtained by integrating the Schechter functions that best reproduce the total IR LF down to $\log(L/L_{\odot})=8$) to $z=4$. The results of integrating the best-fitting curve for our observed total IR LF in each z -bin are shown as black filled circles (the grey filled area is the $\pm 1\sigma$ uncertainty locus) and compared with estimates from previous mid-IR surveys (magenta filled area, Le Floc’h et al. 2005; orange filled triangles, Caputi et al. 2007; blue open triangles, Rodighiero et al. 2010a; and green open circles, Magnelli et al. 2011). The upward pointing arrow in the highest- z bin means that, due to the large fraction of photometric redshifts and the fact that the PEP selection might miss high- z sources, our $3.0 < z < 4.2$ ρ_{IR} estimate is likely to be a lower limit.

ever, that since this population is never dominant in our IR survey, we do not expect that contamination related to accretion activity occurring in these objects (mainly at high- z) can significantly affect the results in terms of ρ_{SFR} .

In Fig. 17 we show ρ_{IR} estimated from our total IR LF and compare it with results obtained from previous IR surveys (Le Floc’h et al. 2005; Caputi et al. 2007; Rodighiero et al. 2010a; Magnelli et al. 2011). In the common redshift intervals ($0 \lesssim z \lesssim 2\text{--}2.5$), we find very close agreement with previous results based on IR data, especially with the Magnelli et al. (2011) derivation. As well as previous findings, ρ_{IR} from PEP shows the rapid rise from $z \sim 0$ to $z \sim 1$, followed by a flattening at higher redshifts. The indications from our survey are that the intermediate redshift flattening is followed by a high redshift decline, which starts around $z \sim 3$. From our data, ρ_{IR} evolves as $(1+z)^{3.0 \pm 0.2}$ up to $z \sim 1.1$, as $(1+z)^{-0.3 \pm 0.1}$ from $z \sim 1.1$ to $z \sim 2.8$, then as $(1+z)^{-6.0 \pm 0.9}$ up to $z \sim 4$.

In the bottom panels of Fig. 18 we plot the different contributions to ρ_{IR} from the different SED populations (left), from the on- and off-MS sources (middle) and from the different mass intervals. We notice a predominance of **spiral**–SED galaxies only at low redshifts ($z < 0.5\text{--}0.6$), when **SF-AGN** begin to dominate ρ_{IR} up to $z \sim 2.5$. The **starburst** SED galaxies are never the prevalent population, although their contribution to ρ_{IR} increases rapidly from the local Universe to $z \sim 1$, then keeps nearly constant to $z \sim 2.5$, to decrease at higher redshifts. The **SF-AGN(SB)** and **SF-AGN(Spiral)** contributions to ρ_{IR} show opposite trends, with the former sharply increasing towards the higher redshifts (dominating at $z > 2$), and the latter prevailing between $z \sim 1$ and $z \sim 2$, then dropping at higher redshifts. **AGN1** and **AGN2** start dom-

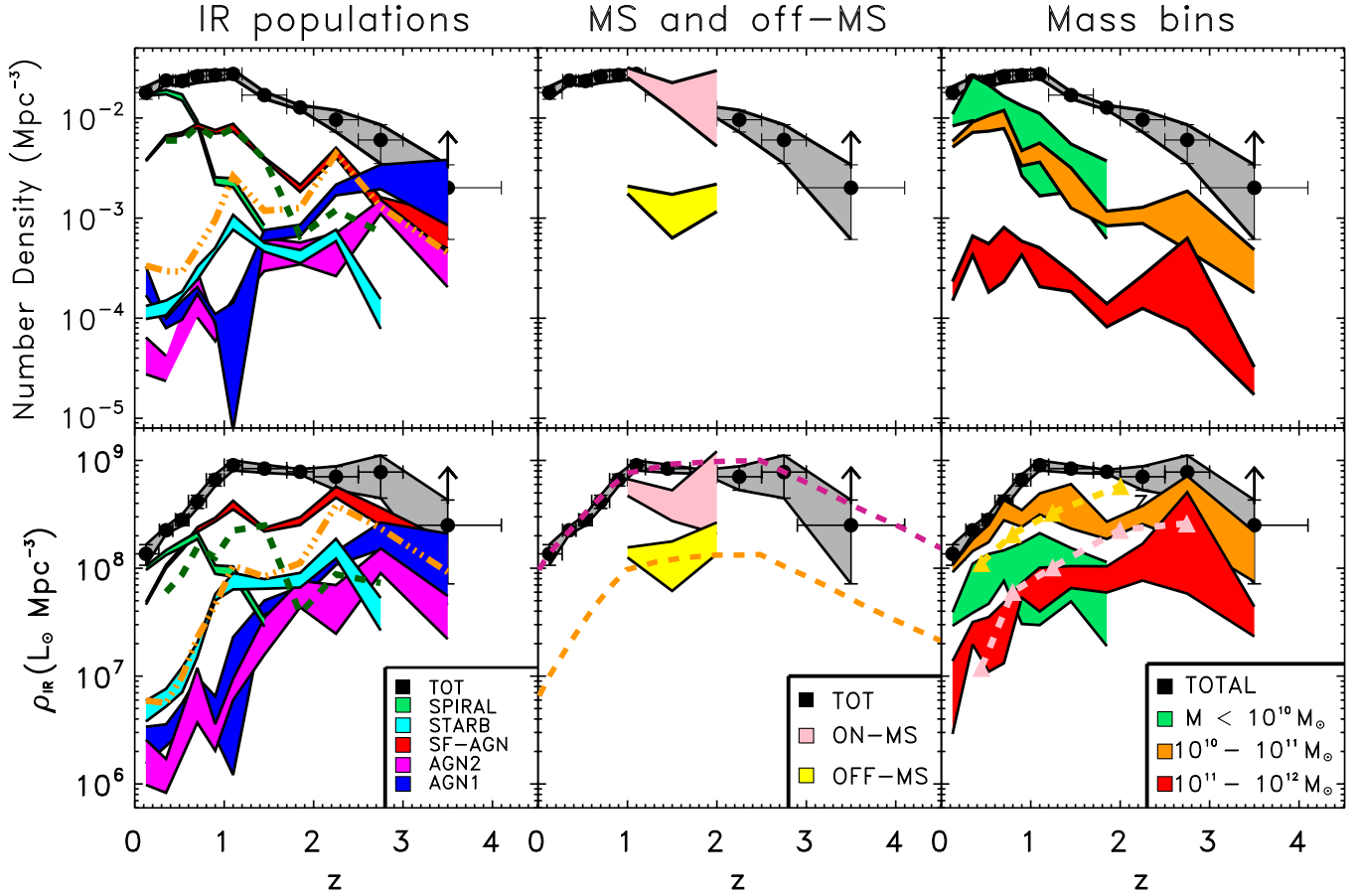


Figure 18. *Top:* Evolution of the comoving number density of PEP sources up to $z \sim 4$ (black filled circles with error-bars within the $\pm 1\sigma$ uncertainty region, represented by the grey filled area). The upward pointing arrow in the highest- z bin means that our $3.0 < z < 4.2$ estimates are likely to be lower limits. *Bottom:* Redshift evolution of the total IR luminosity density to $z=4$. To compute the number and IR luminosity density, we integrate the Schechter functions that best reproduce the different populations/mass/sSFR-classes, down to $\log(L/L_\odot)=8$. The black filled circles and the grey dashed area in all the three panels represent our PEP derived ρ_{IR} and its $\pm 1\sigma$ uncertainty region, as shown in Fig. 17. In the *left* panels we show the number (*top*) and luminosity density (*bottom*) of the different IR populations (green filled area, *spiral*; cyan, *starburst*; red, *SF-AGN*; magenta, *AGN2*; and blue, *AGN1*). The contribution of *SF-AGN* sources, sub-divided on the basis of their SED resemblance to *spiral* or *starburst* templates, are shown by the dark-green dashed (*SF-AGN(Spiral)*) and orange dot-dot-dot-dashed (*SF-AGN(SB)*) lines, respectively. In the *middle* panels we show the uncertainty regions of the relative contribution in number and luminosity density of the sources on- and off- the SFR–stellar mass MS (Elbaz et al. 2007; Daddi et al. 2007), as pink and yellow filled areas, respectively. Our derivations have been compared to those of Bethermin et al. (2012) for on- and off-MS sources (converted from ρ_{SFR} to ρ_{IR} using the Kennicutt 1998 relation), which are represented by the purple and orange dashed lines respectively. In the *right* panels we show the relative contribution to the number and luminosity density of sources with different masses (green, $8.5 < \log(M/M_\odot) < 10$; orange, $10 < \log(M/M_\odot) < 11$; and red, $11 < \log(M/M_\odot) < 12$). For comparison, in the *bottom right* panel we plot also the results of Santini et al. (2009) in the GOODS-S field in similar mass intervals (light-orange triangles and dashed line: $9.77 < \log(M/M_\odot) < 10.77$; pink triangles and dashed line: $\log(M/M_\odot) > 10.77$).

inating the IR luminosity density at $z \gtrsim 2.5$, with their ρ_{IR} always rising from $z \sim 0$, then remaining almost constant (or slightly decreasing) towards the higher redshifts. Note how these two populations of AGN evolve similarly, as an indication of their same intrinsic nature and of the absence of any significant bias in the far-IR selection.

The contributions to ρ_{IR} of MS and off-MS sources (pink and yellow filled regions, respectively, in the *bottom middle* panel of Fig. 18) stay nearly constant between $z \sim 0.8$ and $z \sim 2.2$. In particular, the off-MS sources contribution stays around 20 per cent of the total ρ_{IR} over the whole $0.8 < z < 2.2$ redshift interval, showing no significant signs of increase (or decrease). Our results strengthen

the role of MS sources (pink shaded regions) in the build up of the stellar mass in galaxies, at all cosmic epochs, with evidence that their role is even increasing from $z \sim 2$ to $z \sim 1$: their number density changes by a factor of ~ 2 , while their luminosity/SFR density remains nearly constant. The importance of the off-MS sources does not show any significant signs of decreasing at lower z , their relative (with respect to the total) number (luminosity) density passing from ~ 9 per cent (22 per cent) at $z \sim 2$ to ~ 6 per cent (19 per cent) at $z \sim 1$. These fractions are relatively different from those found by Rodighiero et al. (2011) (off-MS galaxies represent only 2 per cent of mass-selected star-forming galaxies and account for only 10 per cent of the cosmic SFR density at $z \sim 2$).

However, we must note that Rodighiero et al. (2011) for their analysis combined far-IR-selected (i.e., SFR-selected) and near-IR-selected (i.e., M_* -selected) star-forming samples, well defining the main sequence, while with our data (SFR-selected only), we are not able to observe any correlation between SFR and stellar mass (see Fig. 15) and barely detect MS objects at $z \sim 2$.

Our results have been compared to the SFR densities (converted to ρ_{IR} using the Kennicutt 1998 relation) for on- and off-MS sources based on the Sargent et al. (2012) model recently derived by Bethermin et al. (2012). These are shown in Fig. 18 as purple and orange dashed lines respectively. The predicted off-MS ρ_{SFR} agrees well with our estimate in the common redshift range, while the predicted one for MS sources is higher than that derived from our data (especially at $z \sim 1.5$). As already discussed regarding the comparison with the Sargent et al. (2012) model, this discrepancy is likely to be ascribed to our difficulty to extrapolate the MS to low SFR values and to the different selection criteria used to separate MS from off-MS sources.

In the bottom right panels of Fig. 18, we show the contribution of the different mass populations to the luminosity density as a function of redshift. Although we detect a similar steep increase of ρ_{IR} versus redshift at $z \lesssim 1$ in both low and high mass galaxies, the evolution in ρ_{IR} of galaxies with different masses is very different, reflecting the downsizing scenario, with ρ_{IR} peaking at higher redshift with increasing mass. Indeed, the IR luminosity density of intermediate-mass objects ($\log(M/M_\odot) = 10-11$) always dominates, increasing up to $z \sim 1$, then remaining nearly constant at higher redshifts (at least up to $z \sim 2.8$). The IR luminosity density of most massive objects increases even more rapidly with redshift (at $z=2$ it was ~ 5 times higher than today) and continues to grow up to $z=3$, where their contribution to ρ_{IR} is ~ 30 per cent of the total and close to that of intermediate mass objects (which contribute ~ 60 per cent at $z > 3$). We compare our results with those of Santini et al. (2009), plotted in the figure as thick dashed lines (light orange: $9.77 < \log(M/M_\odot) < 10.77$; pink: $\log(M/M_\odot) > 10.77$; with a Chabrier IMF used to determine our masses), showing very similar trends and values for both intermediate- and higher-mass galaxies. Our analysis of high mass galaxies extends up to $z \sim 4$, finding that for the most massive galaxies ρ_{IR} continues to rise even at $z \geq 2$, with an apparent peak at $z = 3$. This result confirms that the formation epoch of galaxies proceeded from high- to low-mass systems.

The values of ρ_{IR} in the different redshift intervals, either the total ones or the contributions from the different classes (SED, mass, sSFR), are reported in Table 9.

5 DISCUSSION

In the previous sections we have discussed the different evolutionary behaviour of different classes of sources, either divided by SED-type, mass or sSFR. In this section we will try to understand “who is who”, discussing which populations are mainly on- or off-MS, which have the larger (smaller) masses, and how and if the relative contributions of these populations vary with redshift.

From Fig. 15 we note that the off-MS sources are dominated by galaxies with AGN-type SEDs. In the lower red-

shift bin considered, MS sources are mainly **spirals** and **SF-AGN(Spiral)**, with a small tail of these populations contributing also to the off-MS. **AGN1** and **AGN2** are prevalently 0.6 dex above the MS (off-MS). Although the mass estimate of **AGN1** suffers of the largest uncertainties, our result suggests that the off-MS population is largely constituted by sources containing an AGN, with a higher fraction of AGN-dominated objects at higher z . This should indicate that the major merger episode likely associated with what is considered a different mode of star-formation (e.g. Wuyts et al. 2011 finds that off-MS galaxies are very compact from their Sérsic index, and quite likely mergers), could also trigger an intense AGN activity, whose presence strongly influences the SFR within the host galaxy. While AGN show preferentially high sSFRs, in our sample they do not seem to prefer higher-mass systems (see Fig. 19). However, this may again be due to our SFR selection, so we miss normal main sequence galaxies especially at low mass (cf. Rodighiero et al. 2011). Therefore, our AGN are still in an intense SF phase, where they are still increasing their stellar mass by actively forming stars. Due to the far-IR selection, in fact, we miss the further phase of quiescent, more massive population, with the SF totally quenched and all the stellar mass already formed and in place. Moreover, our result is not in conflict with the consensus in the literature (i.e. that X-ray AGN above a certain X-ray luminosity do prefer massive galaxies), since our classification is not in AGN luminosity cut, but more a cut in $L_{\text{AGN}}/L_{\text{SF}}$ ratio. The recent result of Mullaney et al. (2012), that L_X/L_{IR} (which is a proxy of $L_{\text{AGN}}/L_{\text{SF}}$) is on average almost independent of stellar mass, implies that our selected AGN are not necessarily hosted by very massive galaxies, as indeed we find.

Our results seem to confirm those of Santini et al. (2012), finding evidence of a higher average SF activity in AGN hosts with respect to inactive galaxies and a more pronounced level of SF enhancement in the hosts of luminous AGN (i.e. our **AGN1** and **AGN2**). Most of the **starburst** galaxies (i.e. with SEDs fitted by local “starburst” templates) are classified on-MS at any redshifts, though they are always in the region between the MS and the threshold. The **starburst** population seems to enhance its sSFR from $z \sim 2$ to $z \sim 1$, occupying the region around the MS at higher z and shifting to higher SFRs (peaking at $\sim 2 \times$ the MS) at lower z . The bulk of the **spiral** population remains around the MS at all redshifts, although they almost disappear from our sample at $z \sim 2$. From Fig. 15 we also note that, while the **SF-AGN** population as a whole occupies both the loci of MS and off-MS sources, when divided into **SF-AGN(SB)** and **SF-AGN(Spiral)**, it shows a very clear segregation (suggesting a different mode of star-formation for the two subclasses). In fact, **SF-AGN(SB)** are mainly concentrated in the off-MS region, while the bulk of the **SF-AGN(Spiral)** is on-MS.

If we mix together all the “ingredients” presented above, we can try to give a global interpretation to our results in terms of galaxy evolution. In agreement with other *Herschel* findings (i.e. Shao et al. 2010; Lapi et al. 2011; Santini et al. 2012; Rosario et al. 2012; Page et al. 2012; Mullaney et al. 2012), we propose the following twofold evolutionary scenario for galaxies and AGN (sketched in a cartoon in Fig. 20):

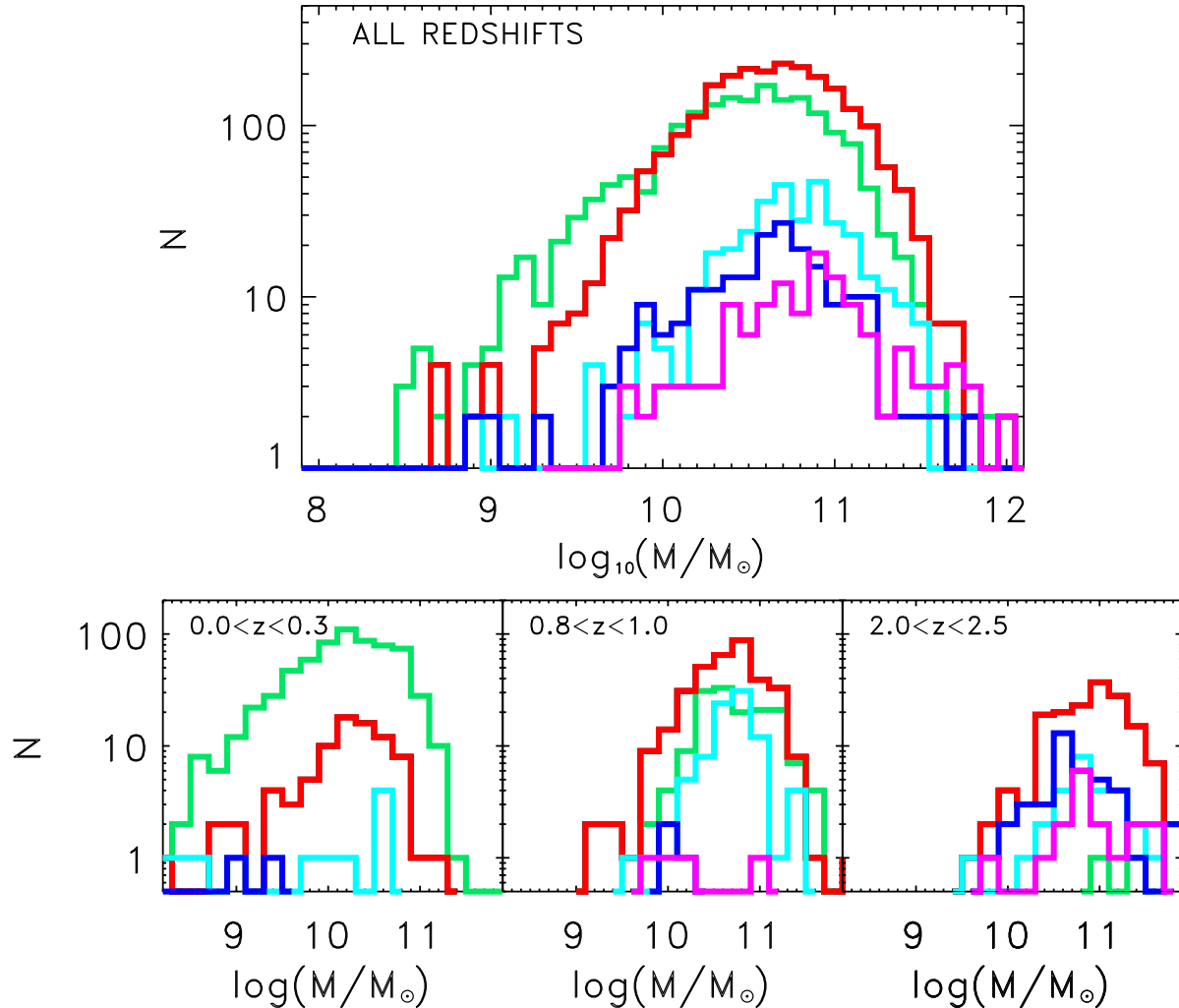


Figure 19. Mass distribution of the PEP 160- μm sources in all the four PEP fields, with different colours corresponding to different SED-classes (green, *spiral*; cyan, *starburst*; red, *SF-AGN*; magenta, *AGN2*; and blue, *AGN1*). The *top* panel shows the mass distribution of the different IR populations at all redshifts, while the *bottom* panels report the mass distribution in three representative redshift bins ($0 < z < 0.3$, $0.8 < z < 1.0$ and $2.0 < z < 2.5$).

• On the one hand, we observe that the sources with AGN-dominated SEDs (either *AGN1* or *AGN2*) and those with a starburst-like SED, but containing a non-dominant AGN (*SF-AGN(SB)*) have a peak in number and luminosity density at $z \sim 2-2.5$, dominating at higher redshifts and rapidly decreasing at lower z . The evolution of *AGN1* and *AGN2*, both in luminosity and in density, is very similar, suggesting the same nature for type 1 and 2 AGN, and the unbiased power of observations in the far-IR band towards orientation/obscuration (affecting both optical and X-ray observations). While dominating the mid-IR part of the SED of these objects, the AGN is not able to explain the high observed far-IR emission, which is mostly powered by star-formation (as for the *SF-AGN(SB)* population, where the AGN never dominates the energetic output, probably due to dust-obscuration). The hosts of these AGN appear to form stars in a very efficient way, placing a large fraction of them above the known stellar mass-SFR MS (see Fig. 15). *AGN1*, *AGN2* and *SF-AGN(SB)* are likely the progenitors of the elliptical galaxies observed nowadays in optical and near-IR

surveys, forming through an intense burst of star-formation occurring during major mergers or in dense nuclear star-forming regions (Granato et al. 2001; Daddi et al. 2010; Wuyts et al. 2011), then followed by a phase of nuclear activity during which their SMBHs grow (i.e. Hopkins et al. 2008a, 2008b; Lapi et al. 2011). It is generally agreed that the SMBHs and their host galaxies are tightly related, with major-mergers being considered the likely process responsible for transporting large amount of gas towards the centre of the merging system, feeding the SMBH and triggering the SF activity. After the intense starburst phase, the AGN are believed to suppress the SF (i.e. Di Matteo et al. 2005), so that the remnant quickly evolves to a red massive spheroid. This picture is strongly supported by the recent *Herschel* results from PEP (Rosario et al. 2012), HERMES (Page et al. 2012; Harrison et al. 2012) and H-ATLAS (Lapi et al. 2011) Surveys. The latter work, in particular, has shown that the bright-end of the IR LFs and counts at high redshift (>1.5) are consistent with the picture (e.g. Granato et al. 2001) predicting the presence of a popula-

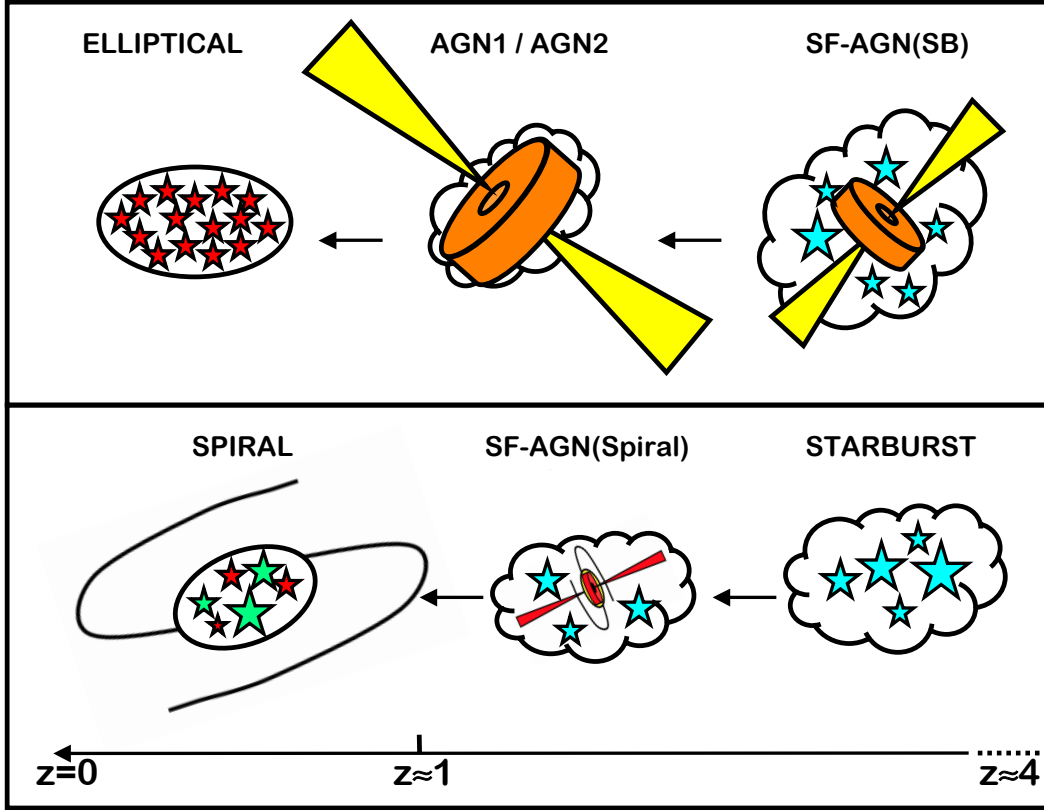


Figure 20. A cartoon showing a possible evolutionary scenario involving AGN and star-forming galaxies and leading to the formation, on the one hand, of local elliptical galaxies, on the other hand, of local spiral galaxies. *Top:* a strong starburst with a growing BH inside (SF-AGN(SB)) transforms into an AGN (either AGN1 or AGN2, depending on orientation) when the BH mass and AGN luminosity have grown enough and the star formation is likely to be quenched by feedback processes. Thereafter the galaxy evolves passively toward a local elliptical. *Bottom:* the initial moderate starburst (starburst, lasting a typical time of $\sim \text{few } 10^8$ yrs) transforms in a less intense starburst also fueling a low-luminosity AGN (SF-AGN(Spiral)), which starts to reveal itself when the starburst activity is fading. Once the AGN is triggered, it heats the dust resulting in an increase in the mid-infrared luminosity (flattening of the SED between 3 and $8 \mu\text{m}$). The SF-AGN(Spiral) system could last about $7\times$ longer than the pure starburst phase before becoming a steady spiral galaxy (spiral).

tion of strongly-obscured, star-forming galaxies with SED appreciably different from those of the local starbursts. X-ray and mid-IR spectroscopic observations (e.g. Alexander et al. 2005, Vliante et al. 2007) of sub-mm selected sources have revealed in many of them the presence of a growing SMBH, powering an obscured AGNs. In this framework, the most likely evolutionary path envisages first a SF-AGN(SB) phase (star-forming galaxy with a growing SMBH inside), then an AGN-dominated phase, and finally the formation of an elliptical galaxy in passive evolution.

- On the other hand, we observe that a significant fraction of our IR selected sources is constituted by moderately star-forming galaxies characterised by an SED similar to that of spiral galaxies, but also suggesting the presence of a low-luminosity AGN (SF-AGN(Spiral)), best-fitted by local Seyfert 1.8/2 templates. The bulk of these objects and their principal contribution to ρ_{IR} are at intermediate redshifts ($1 \lesssim z \lesssim 2$), while they decrease between $z \sim 1$ and $z = 0$, as the spiral population rises. Most of the SF-AGN(Spiral)-, spiral- and starburst-SED galaxies occupy the region

around the SFR-stellar mass MS at any redshifts, suggesting a steady mode of star-formation rather than a “starburst” one for these three populations, whose evolution is likely connected. Indeed, Mullaney et al. (2012) have recently shown that at least up to $z \sim 2$, SMBHs have grown together with their host galaxies in star-forming galaxies, irrespective of host galaxy mass and triggering mechanism. Given the number densities (Fig. 18), evolutionary trends (Fig. 11), SFRs and masses (Figs. 15 and 19) of the SF-AGN(Spiral), starburst and spiral populations, we suggest that these three classes of objects might constitute different phases in the life of a galaxy undergoing secular evolution. The gas in moderate starburst galaxies, undergoing a burst of enhanced SF due either to gravitational interactions or disk instabilities (typical burst duration time of the order of a few 10^8 yr), might also fuel a low-luminosity AGN, which starts to reveal itself when the starburst activity is fading. Given the relatively high stellar masses found for the bulk of the SF-AGN(Spiral) ($\log(M/M_\odot) \sim 10-11$, see Fig. 15), and the almost constant M_{BH}/M_\star ratio (of $\sim 1-2 \times 10^{-3}$) recently

suggested by Mullaney et al. (2012) for all the $0.5 < z < 2.5$ star-forming galaxies, they are likely to contain relatively massive BHs ($M_{\text{BH}} \sim 10^7 - 10^8 M_{\odot}$). They can therefore have either low values of their radiative efficiency or low values of their accretion mass rate \dot{m} ($\dot{m} = \dot{M}/\dot{M}_{\text{Edd}} < 0.01$). AGN with low radiative efficiency or low accretion mass rates are generally called radiatively inefficient accretion flows (RIAFs – Narayan 1995–, which include also the advection-dominated accretion flow, ADAF – Narayan 1994). Low \dot{m} AGN are more difficult to detect, since they often are less luminous than their hosts. Because of this, the nuclear emission is diluted by the host galaxy’s emission and many of them are likely to be classified as “normal galaxies” in most surveys if the AGN luminosity is less than that of the host (e.g. Hopkins et al. 2009). Given the relative number densities of the **starburst** and **SF-AGN(Spiral)** populations at $1 < z < 2$, we hypothesize a typical duration of the **SF-AGN(Spiral)** phase about 7 times longer than the **starburst** one (typical burst duration $\sim 10^8$ yr). Then, after a typical time of $\sim 7 \times 10^8$ yr, the AGN activity stops and these objects, whose number density decreases at $z \lesssim 1$, are likely to become steady **spiral** galaxies (rapidly increasing between $z \sim 1$ and $z = 0$) at lower redshifts.

6 CONCLUSIONS

We have used the 70-, 100-, 160-, 250-, 350- and 500- μm data from the cosmological guaranteed time *Herschel* surveys, PEP and HerMES, in the GOODS-S and -N, ECDFS and COSMOS, to characterise the evolution of the IR luminosity function and luminosity density of PACS selected sources across the redshift range $0 \lesssim z \lesssim 4$. Evolution is well constrained by our data up to $z \sim 3$, strong hints of evolution are derived at $3 < z < 4$. In the present work we have:

(i) completely characterised the multi-wavelength SEDs of the PEP sources by performing a detailed SED-fitting analysis and comparison with known template library of IR populations. Sources have been classified, based on their broad-band SEDs, in five main classes: **spiral**, **starburst**, **SF-AGN**, **AGN1** and **AGN2**.

(ii) computed the rest-frame LFs at 35, 60 and 90 μm up to $z \sim 4$ from the 70-, 100- and 160- μm selected samples respectively.

(iii) integrated the SEDs over $\lambda_{\text{rest}} = 8 - 1000 \mu\text{m}$ and computed the total IR LF up to $z \sim 4$ and studied its evolution with redshift, finding strong luminosity evolution $\propto (1+z)^{3.55 \pm 0.10}$ up to $z \sim 1.85$, and $\propto (1+z)^{1.62 \pm 0.51}$ between $z \sim 1.85$ and $z \sim 4$, combined with a negative density evolution $\propto (1+z)^{-0.57 \pm 0.22}$ up to $z \sim 1.1$ and $\propto (1+z)^{-3.92 \pm 0.34}$ at $z > 1.1$ and up to $z \sim 4$.

(iv) derived the evolution of the comoving total IR luminosity density, which is found to increase as $(1+z)^{3.0 \pm 0.2}$ up to $z \sim 1.1$, then to remain nearly constant (decrease as $(1+z)^{-0.3 \pm 0.1}$) from $z \sim 1.1$ to $z \sim 2.8$, and to decrease as $(1+z)^{-6.0 \pm 0.9}$ up to $z \sim 4$.

(v) found that the evolution derived for the global IR LF is indeed a combination of different evolutionary paths: the IR population does not evolve all together “as a whole”, as is often assumed in the literature, but is composed of different galaxy classes evolving differently: the **spiral**-SED galaxies dominate ρ_{IR} only at low redshifts ($z \lesssim 0.5 - 0.6$), then **SF-AGN**

dominate up to $z \sim 2.5$, while **AGN1** and **AGN2** start dominating the IR luminosity density only at $z \gtrsim 2.5$.

(vi) derived the relative contribution to ρ_{IR} of MS and off-MS sources, which keep nearly constant between $z \sim 0.8$ and $z \sim 2.2$, with the MS population always dominating. The contribution to ρ_{IR} of the off-MS sources shows no significant signs of increase with z (from ~ 19 per cent at $z \sim 0.8 - 1.25$ to $\gtrsim 22$ per cent at $z \sim 1.8 - 2.2$).

(vii) derived very different evolutionary behaviour in terms of different contributions to ρ_{IR} , for galaxies with different masses, reflecting the downsizing scenario (ρ_{IR} peaks at higher redshift with increasing mass). Intermediate-mass objects ($\log(M/M_{\odot}) = 10 - 11$) always dominate the IR luminosity density, increasing with redshift up to $z \sim 1$, then remaining nearly constant at higher redshifts (at least up to $z \sim 2.8$), while the contribution of most massive objects increases even more rapidly with z (at $z \sim 2$ it was ~ 5 times higher than today) and continues to grow up to $z \sim 3$.

(viii) described a possible twofold evolutionary scenario for IR sources: on the one hand, **AGN1** and **AGN2**, representing the same population, after an intense starburst phase (due to a major merging event, appearing as **SF-AGN(SB)** SED galaxies), suppress the SF (and shine as AGN) and evolve to red massive spheroids; on the other hand, the **SF-AGN(Spiral)** galaxies represent a phase in the life of a star-forming galaxy, following a moderate burst of SF (**starburst**, with SEDs like those of local starburst galaxies) and preceding the formation of a steady **spiral** galaxy as we observed in the local Universe.

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Table 3: Rest-frame 35 μm luminosity function

$\log(L_{35}/L_{\odot})$	from the GOODS-S 70- μm sample						
	$\log(\Phi/\text{Mpc}^{-3} \text{ dex}^{-1})$						
	$0.0 \leq z < 0.4$	$0.4 \leq z < 0.8$	$0.8 \leq z < 1.2$	$1.2 \leq z < 1.8$	$1.8 \leq z < 2.5$	$2.5 \leq z < 3.5$	$3.5 \leq z < 4.5$
7.9–8.3	-1.34 ± 0.29						
8.3–8.7	-1.62 ± 0.04						
8.7–9.1	-2.01 ± 0.20						
9.1–9.5	-2.38 ± 0.14						
9.5–9.9	-2.63 ± 0.12	-2.26 ± 0.29					
9.9–10.3	-2.96 ± 0.15	-2.81 ± 0.14					
10.3–10.7	-3.11 ± 0.18	-2.70 ± 0.06					
10.7–11.1	-3.50 ± 0.43	-3.23 ± 0.09	-3.03 ± 0.09				
11.1–11.5	-3.90 ± 0.43	-3.98 ± 0.22	-3.50 ± 0.10	-3.19 ± 0.15			
11.5–11.9		-4.59 ± 0.43	-4.35 ± 0.25	-4.14 ± 0.17	-3.82 ± 0.20		
11.9–12.3				-4.41 ± 0.19	-4.14 ± 0.19		
12.3–12.7					-4.68 ± 0.32	-4.27 ± 0.30	
12.7–13.1						-5.27 ± 0.43	-5.13 ± 0.43
13.1–13.5							-5.34 ± 0.43
from the combined 100- μm sample							
7.5–7.9	-1.74 ± 0.43						
7.9–8.3	-1.64 ± 0.23						
8.3–8.7	-2.00 ± 0.19						
8.7–9.1	-2.01 ± 0.09						
9.1–9.5	-2.37 ± 0.03	-2.43 ± 0.37					
9.5–9.9	-2.53 ± 0.02	-2.48 ± 0.19					
9.9–10.3	-2.73 ± 0.02	-2.66 ± 0.08					
10.3–10.7	-3.14 ± 0.03	-2.56 ± 0.03	-2.79 ± 0.10				
10.7–11.1	-3.91 ± 0.07	-3.34 ± 0.02	-2.91 ± 0.06				
11.1–11.5	-4.73 ± 0.18	-4.03 ± 0.04	-3.46 ± 0.03	-3.13 ± 0.07	-3.24 ± 0.32		
11.5–11.9	-5.51 ± 0.43	-4.90 ± 0.10	-4.09 ± 0.03	-3.77 ± 0.04	-3.38 ± 0.10		
11.9–12.3		-5.60 ± 0.22	-4.90 ± 0.07	-4.37 ± 0.03	-3.98 ± 0.07	-3.18 ± 0.29	
12.3–12.7			-6.16 ± 0.31	-5.34 ± 0.09	-4.54 ± 0.04	-4.30 ± 0.12	
12.7–13.1				-6.46 ± 0.31	-5.48 ± 0.09	-4.99 ± 0.07	-4.40 ± 0.39
13.1–13.5						-6.03 ± 0.15	-5.91 ± 0.26
13.5–13.9						-7.04 ± 0.43	

Table 4: Combined rest-frame 60 μm luminosity function

$\log(L_{60}/L_{\odot})$	$\log(\Phi/\text{Mpc}^{-3} \text{ dex}^{-1})$						
	$0.0 \leq z < 0.4$	$0.4 \leq z < 0.8$	$0.8 \leq z < 1.2$	$1.2 \leq z < 1.8$	$1.8 \leq z < 2.5$	$2.5 \leq z < 3.5$	$3.5 \leq z < 4.5$
7.5–7.9	-1.37 ± 0.43						
7.9–8.3	-1.54 ± 0.32						
8.3–8.7	-1.82 ± 0.20						
8.7–9.1	-1.75 ± 0.09						
9.1–9.5	-2.22 ± 0.04						
9.5–9.9	-2.36 ± 0.03	-2.25 ± 0.27					
9.9–10.3	-2.58 ± 0.02	-2.70 ± 0.12					
10.3–10.7	-2.77 ± 0.02	-2.46 ± 0.04					
10.7–11.1	-3.28 ± 0.03	-3.00 ± 0.02	-2.70 ± 0.07				
11.1–11.5	-4.13 ± 0.09	-3.41 ± 0.02	-3.12 ± 0.04	-3.09 ± 0.11			
11.5–11.9	-4.91 ± 0.22	-4.26 ± 0.05	-3.61 ± 0.02	-3.47 ± 0.06	-3.08 ± 0.22		
11.9–12.3		-5.03 ± 0.11	-4.31 ± 0.04	-3.95 ± 0.03	-3.63 ± 0.10	-3.22 ± 0.31	
12.3–12.7		-6.20 ± 0.43	-5.31 ± 0.12	-4.73 ± 0.05	-4.22 ± 0.06	-4.09 ± 0.15	
12.7–13.1				-5.98 ± 0.18	-5.08 ± 0.06	-4.76 ± 0.13	-4.45 ± 0.43
13.1–13.5					-6.88 ± 0.43	-5.58 ± 0.22	-5.39 ± 0.43
13.5–13.9						-6.93 ± 0.43	-6.20 ± 0.43

Table 5: Combined rest-frame 90 μm luminosity function

$\log(L_{90}/L_{\odot})$	$\log(\Phi/\text{Mpc}^{-3} \text{ dex}^{-1})$						
	$0.0 \leq z < 0.4$	$0.4 \leq z < 0.8$	$0.8 \leq z < 1.2$	$1.2 \leq z < 1.8$	$1.8 \leq z < 2.5$	$2.5 \leq z < 3.5$	$3.5 \leq z < 4.5$
7.9–8.3	-1.12 ± 0.33						
8.3–8.7	-2.11 ± 0.43						
8.7–9.1	-1.89 ± 0.19						
9.1–9.5	-2.21 ± 0.07						
9.5–9.9	-2.31 ± 0.04						
9.9–10.3	-2.51 ± 0.03	-2.06 ± 0.24					
10.3–10.7	-2.69 ± 0.02	-2.57 ± 0.05					
10.7–11.1	-3.04 ± 0.03	-2.98 ± 0.02	-2.97 ± 0.08	-2.90 ± 0.20			
11.1–11.5	-3.91 ± 0.07	-3.31 ± 0.02	-3.02 ± 0.03	-3.21 ± 0.13			
11.5–11.9	-4.61 ± 0.15	-4.10 ± 0.04	-3.65 ± 0.02	-3.58 ± 0.06	-3.47 ± 0.11	-3.84 ± 0.30	
11.9–12.3		-5.06 ± 0.12	-4.37 ± 0.04	-4.12 ± 0.03	-3.97 ± 0.08	-4.08 ± 0.16	-4.69 ± 0.22
12.3–12.7			-5.76 ± 0.19	-5.02 ± 0.06	-4.26 ± 0.10	-4.48 ± 0.10	-5.36 ± 0.41
12.7–13.1				-6.47 ± 0.31	-5.50 ± 0.10	-5.13 ± 0.11	-5.86 ± 0.19
13.1–13.5						-6.61 ± 0.43	

Table 6: PEP total IR luminosity function

$\log(L_{\text{IR}}/L_{\odot})$	$\log(\Phi/\text{Mpc}^{-3} \text{ dex}^{-1})$										
	$0.0 \leq z < 0.3$	$0.3 \leq z < 0.45$	$0.45 \leq z < 0.6$	$0.6 \leq z < 0.8$	$0.8 \leq z < 1.0$	$1.0 \leq z < 1.2$	$1.2 \leq z < 1.7$	$1.7 \leq z < 2.0$	$2.0 \leq z < 2.5$	$2.5 \leq z < 3.0$	$3.0 \leq z < 4.2$
8.5–9.0	-2.21±0.43										
9.0–9.5	-2.18±0.09										
9.5–10.0	-2.28±0.04										
10.0–10.5	-2.50±0.03	-2.37±0.11									
10.5–11.0	-2.71±0.02	-2.64±0.04	-2.61±0.08	-2.27±0.15							
11.0–11.5	-3.49±0.06	-3.12±0.03	-2.99±0.04	-2.89±0.05	-3.09±0.08	-2.80±0.09	-2.93±0.18				
11.5–12.0	-4.79±0.25	-4.29±0.10	-3.89±0.05	-3.53±0.03	-3.24±0.04	-3.17±0.06	-3.29±0.06	-3.76±0.13			
12.0–12.5	-5.27±0.43	-5.58±0.43	-5.50±0.31	-4.75±0.09	-4.23±0.05	-4.00±0.03	-3.81±0.03	-3.96±0.11	-3.53±0.08	-3.75±0.21	
12.5–13.0				-5.79±0.31	-5.74±0.25	-5.18±0.12	-4.85±0.05	-4.42±0.04	-4.40±0.04	-4.15±0.11	-4.65±0.14
13.0–13.5							-6.48±0.31	-6.01±0.22	-5.79±0.13	-5.11±0.07	-5.75±0.13
13.5–14.0										-6.54±0.31	-7.18±0.43

Table 9: PEP total IR luminosity density

	$\rho_{\text{IR}}/(10^8 L_{\odot} \text{ Mpc}^{-3})$										
	$0.0 \leq z < 0.3$	$0.3 \leq z < 0.45$	$0.45 \leq z < 0.6$	$0.6 \leq z < 0.8$	$0.8 \leq z < 1.0$	$1.0 \leq z < 1.2$	$1.2 \leq z < 1.7$	$1.7 \leq z < 2.0$	$2.0 \leq z < 2.5$	$2.5 \leq z < 3.0$	$3.0 \leq z < 4.2$
TOTAL	1.36±0.29	2.25±0.18	2.82±0.32	4.17±0.58	6.61±0.83	9.02±0.95	8.39±0.73	7.85±0.47	7.06±1.75	7.80±3.33	2.50±1.78
spiral	1.00±0.04	1.39±0.07	1.68±0.12	2.02±0.15	0.98±0.09	0.94±0.10	0.32±0.03				
starburst	0.05±0.01	0.06±0.01	0.10±0.02	0.18±0.03	0.54±0.04	0.77±0.13	0.72±0.07	0.78±0.12	1.55±0.32	0.40±0.13	
SF-AGN	0.47±0.01	1.03±0.04	1.52±0.06	2.28±0.12	2.82±0.12	3.87±0.36	2.28±0.09	2.75±0.23	5.11±0.60	3.07±0.58	1.29±0.49
AGN2	0.02±0.01	0.012±0.004		0.08±0.04	0.03±0.01	0.08±0.02	0.25±0.10	0.60±0.17	0.47±0.22	1.21±0.31	0.39±0.17
AGN1	0.025±0.009	0.029±0.007	0.04±0.01	0.07±0.02	0.05±0.02	0.12±0.10	0.44±0.06	0.59±0.08	1.25±0.18	2.07±0.59	1.28±0.82
SF-AGN(SB)	0.06±0.04	0.06±0.01	0.10±0.01	0.23±0.03	0.46±0.05	1.05±0.15	0.84±0.06	1.12±0.12	3.83±0.74	2.35±0.62	0.92±0.75
SF-AGN(Spiral)	0.40±0.02	0.60±0.05	0.87±0.01	1.56±0.18	1.26±0.16	2.23±0.38	2.49±0.29	0.38±0.08	0.88±0.55	0.73±0.32	
log(M/M _⊙)=8.5–10	0.34±0.05	0.74±0.36	0.91±0.44	1.11±0.35	1.01±0.71	1.21±0.91	1.02±0.52				
log(M/M _⊙)=10–11	0.99±0.08	1.65±0.21	2.14±0.40	3.54±0.75	2.94±0.52	4.05±0.90	4.18±1.85	2.28±0.41	3.18±0.61	4.47±2.69	1.48±0.71
log(M/M _⊙)=11–12	0.08±0.05	0.26±0.06	0.23±0.12	0.31±0.18	0.63±0.07	0.71±0.31	0.85±0.20	0.82±0.22	1.23±0.46	2.82±2.24	0.34±0.10
	0.8<z<1.25	1.25<z<1.8	1.8<z<2.2								
on-MS	5.71±1.01	4.02±1.27	7.12±5.03								
off-MS	1.41±0.15	1.19±0.58	2.01±0.65								